

Variability in the 2MASS Calibration Fields: A Search for Transient Obscuration Events

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ABSTRACT

We searched the light curves of over 40000 stars in the 2 Micron All Sky Survey (2MASS) calibration database, spanning approximately 4 years, for objects that have significant day long dimming events. We also searched the multi-color light curves for red dimming events that could be due to transient extinction. In the color independent sigma-limited search, we found 46 previously unknown eclipsing binaries, 6 periodic variable stars likely to be intrinsic pulsators and 21 young stellar objects in the ρ Ophiucus star formation region previously studied by Parks et al. 2013. An additional 11 objects exhibited dimming events and most of these are unclassified.

The search for red dimming events primarily reveals a population of low luminosity active galaxies that become bluer when they are brighter, and variable young stellar objects exhibiting high cross-correlation coefficients between color and brightness. The young stellar objects exhibit brightness and color variations in the direction of interstellar extinction whereas the active galaxies can have a bowed distribution in color and magnitude with reduced variation in color when the object is brightest.

Among the objects that are usually quiescent (not strongly variable), we failed to find any dimming events deeper than 0.2 magnitude and lasting longer than a day. Two of the least embedded young stellar objects, however, dimmed by 0.2 mag for longer than a day without strong color variation. Having failed to find new exotic objects, we explored ways to eliminate commonly found objects so that a larger number of objects may be searched. We find that all but one eclipsing binary are excluded by requiring moderate color variation during a dimming event and most of the active galaxies are excluded by placing a limit on the standard deviation of the magnitude distribution in the light curve.

1 INTRODUCTION

One star in a young stellar binary system can host a circumstellar disc and the disc can periodically occult the other star as seen from a distance (e.g., ϵ Aurigae, Guinan & DeWarf 2002; Kloppenborg et al. 2010; Chadima et al. 2011; EE-Cep, Mikolajewski & Graczyk 1999; Graczyk et al. 2003; Mikolajewski et al. 2005; Galan et al. 2012; OGLE-LMC-ECL-17782, Graczyk et al. 2011; OGLE-LMC-ECL-11893, Dong et al. 2014). Occulting disks are valuable as they would allow the structure of a distant disk to be studied on smaller spatial scales than possible with direct imaging and it would be possible to study their composition through absorption of star light. Unfortunately only a few systems are currently known. Because discs can be large, the probability that a

randomly oriented but young system exhibits eclipses may be fairly high (Mamajek et al. 2012; Meng et al. 2014). The scarcity of known occulting disks might then be because there have been few (or no) systematic searches for them. Circumbinary disks can also periodically occult a star (e.g., KH 15D, Kearns & Herbst 1998; Winn et al. 2004; V718 Per, Grinin et al 2008). Periodic and aperiodic occultations have been detected in young stellar objects (Bouvier et al. 2007; Grinin et al 2008; Morales-Calderon et al. 2011; Parks et al. 2013; Wolk et al. 2013; Cody et al. 2014), whereas transient single, long, and deep occultation events of pre-main sequence stars have also been reported (Rodriguez et al. 2013; Mamajek et al. 2012).

The Two Micron All-Sky Survey (Skrutskie et al. 2006;

2MASS) imaged the entire sky in three near infrared wavelength bands, J, H, K_s , between 1997 and 2001. Photometric calibration for 2MASS relied upon repeated observations of 35 calibration fields that are distributed at approximately 2 hour intervals in right ascension near declinations of 0 and ± 30 deg. All source extractions from the calibration scans are available in the 2MASS Calibration Point Source Working Database (Cal-PSWDB). The database contains over 191 million source extractions. Further descriptions of the Cal-PSWDB are available by Cutri et al. (2006), and Plavchan et al. (2008) who found 247 red variable objects, including extragalactic variables and 23 periodic variable stars. Using one of the calibration fields, Parks et al. (2013) recently presented a variability study of the ρ Ophiucus molecular cloud and star formation region, discovering 22 new cluster members via their near-infrared variability and periodicity in 35 young stellar objects.

While numerous photometric monitoring projects search for objects that brighten, such as novae or microlenses, there have been few systematic searches for rare, individual dimming events. Single, isolated, dimming events are often ignored because they can be caused by poor photometric measurements that occur because of seeing variations, scintillation, clouds, imaging artifacts, and source confusion. Variability studies search for periodic phenomena, including planetary transits, (e.g., Collier Cameron et al. 2006; Palaversa et al. 2013), use the moments of the magnitude distribution to find and classify variables (e.g., Ciocca 2013; Graczyk & Eyer 2010) or combine periods found in one band with color or spectral information (not necessarily as a function of time) to classify variable objects (e.g., Palaversa et al. 2013; Masci et al. 2014). Because photometric pipelines can report false dimming events, rare and deep dimming events have only been confirmed when the obscuration is periodic (e.g., the eclipsing disk OGLE-LMC-ECL-11893 with a 468 day period; Dong et al. 2014) or multiple observation platforms observed the same drop in magnitude (e.g., RW Aur A; Rodriguez et al. 2013).

Here we focus on searching for potentially rare events and so we do not restrict our study to periodic stars or stars that are sufficiently variable that they can be found via statistical measurements of the magnitude distribution. We carry out a multi-band search for obscuration events in a fairly large sample, over 40000 stars, the brighter stars in the 2MASS Cal-PSWDB. The Cal-PSWDB is uniformly observed and processed, and has well characterized astrometric and photometric measurements. It is also a multi-band photometric survey with all bands measured at each observation time. As variability can occur on multiple timescales, simultaneous (or near simultaneous) multi-band measurements are needed to disentangle brightness variations from color variations. Variability is often studied in a single photometric wavelength band, however, if obscuration is due to extinction from dust, then the depth of a dimming event or eclipse can depend upon wavelength. For example it is possible to differentiate between star spots and obscuration using near-infrared color (Morales-Calderon et al. 2011; Parks et al. 2013; Wolk et al. 2013; Cody et al. 2014). The eclipses of the OGLE-LMC-ECL-11893 and EE-Cep disk systems are deeper in B band than I band (Galan et al. 2012; Dong et al. 2014). In a three band photometric study, it is also possible to check the uniformity or noise level of measurements

at a particular time by comparing the measurements in the three bands. Here we carry out a search for dimming events in single bands and also search for dimming events that are red and so could be due to transient levels of extinction.

2 SEARCHING FOR TRANSIENT LEVELS OF OBSCURATION

Source lists in each 2MASS calibration tile are available at http://www.ipac.caltech.edu/2mass/releases/allsky/doc/seca7_4.html. Using those source lists we selected all objects with $J < 16.0$ or $K < 15.0$ where J and K are the J - and K -band magnitudes, respectively. This selection ensures bright enough sources that individual photometric measurements are accurate with standard deviation approximately 0.1 mag or better; (this procedure was also followed by Plavchan et al. 2008; see their Figure 7). Light curves for 48,976 objects, satisfying these criteria and sorted in order of time of observation, were downloaded from the 2MASS Cal-PSWDB by querying all observations within 3" of the positions on the sky listed in the source lists¹. Not all of the 48,976 objects gave light curves with many data points, consequently the actual number of stars that can be studied for variability is somewhat smaller (see the discussion by Plavchan et al. 2008). We counted 40,486 objects with greater than 300 observation times (all 3 bands are usually observed at each time) and of these 22,983 objects with greater than 1000 observation times.

From each light curve we measured the mean, μ_J, μ_H, μ_K , and standard deviation, $\sigma_J, \sigma_H, \sigma_K$, of the magnitude distribution in each band. To compute the means and standard deviations, observation points were weighted by the variance of each photometric measurement. We only used light curve points with non *null* listed errors, i.e., those with good (based on 2MASS quality flags) photometric measurements. We also box smoothed the light curves over a 1 day long observation interval by computing the weighted mean and standard deviation of the magnitudes observed within a day of each observation time. Uncertainties for each data point in the box-smoothed curve were calculated from the data points used to make the averages.

We carried out two searches. First in each band, separately, we searched for day long dimming events that are significant compared to the dispersion of points in the light curves. Second, we searched for dimming events that are comparatively red.

For objects found in the searches we investigate a possible correlation between color and brightness. We compute the cross correlation coefficient² using only data points with non-null J and K magnitudes

$$R = \frac{\sum_i (J_i - \mu_J)(J_i - K_i - (\mu_J - \mu_K))}{\sqrt{\sum_i (J_i - \mu_J)^2 \sum_i (J_i - K_i - (\mu_J - \mu_K))^2}} \quad (1)$$

¹ A example query of an object with RA=28.651857° and DEC=+0.225549° (J2000): `wget "http://irsa.ipac.caltech.edu/cgi-bin/Gator/nph-query?outfmt=1&order=jdate&objstr=28.651857+0.225549&spatial=Cone&radius=3&radunits=arcsec&catalog=pt_src_c&selcols=ra,dec,j_m,j_msig,h_m,h_msig,k_m,k_msig,jdate" -O file.tbl`

² <http://mathworld.wolfram.com/CorrelationCoefficient.html>

Here $R = 1$ for a light curve that reddens as it dims (with no scatter or noise) and $R = 0$ if there is no correlation between brightness and color. We also fit (using least mean squares) a line to points on a color/magnitude plot where observations at different times in J band are compared to their $J - K$ color (as done by Wolk et al. 2013). We recorded the slope

$$b \equiv \frac{\Delta J}{\Delta(J - K)} \quad (2)$$

of the best fitting line on the color/magnitude plot. If the color variations and dimming are due to extinction from interstellar dust then we expect a slope $b = 1.6$ (using a standard extinction law, Table 1 Mathis 1990).

3 COLOR INSENSITIVE SEARCH FOR DIMMING EVENTS

We first describe a color insensitive, but sigma-limited, search for dimming events. For each night's observations we measured the median magnitude in each band. We required there to be greater than 4 good (non *null* error) observations during the night. We also measured the median value and standard deviation for the night's observations, μ_n, σ_n . A candidate dimming event required the night's median value to be 1) fainter than 2σ (dispersion of entire light curve) from the mean of the entire light curve, 2) fainter than $1.5\sigma_n$ (night's dispersion) from the mean of the entire light curve and 3) greater than 0.1 mag from the mean of the entire light curve. Here σ is the standard deviation computed from the entire light curve in the band used for the search. This procedure pulls out rare dimming events compared to the distribution of the entire light curve. Each night's observations was tested separately so if a dimming event lasts longer than a night we flag the object each night that the object is dim.

At each candidate dimming event, we visually inspected the light curves. Dimming events in noisy light curves (i.e., objects with large dispersions) were discarded. The noise was usually due to faintness of the star, but was sometimes due to confusion with nearby sources. Care was taken to check the time of individual dimming events. If more than 1 star in the same calibration tile showed a dimming event on the exact same night, the events were discarded. In some calibration tiles, many stars exhibited false dimming events on the same night.

For each star with a candidate dimming event, we searched the VizieR database³ for variability. Excluding the ρ Ophiucus region, we found 15 previously known periodic systems, listed in the American Association of Variable Star Observer's (AAVSO) International Variable Star Index⁴ and one previously known variable extragalactic source (2MASS 14584479+3720216), previously noted by Plavchan et al. (2008). The previously known periodic stars are 2MASS

03315574+3703122 and 2MASS 03314856+3723373 (eclipsing binaries found by J. Greaves⁵), 2MASS J08255405-3908441, 2MASS J18510479-0442005, 2MASS J18512034-0426311, 2MASS J18512261-0409084, 2MASS J18512929-0412407, and 2MASS J19020989-0439440, that were previously noted as eclipsing binaries by Plavchan et al. (2008), and the following that are listed in the AAVSO index: 2MASS J03322127+3701153 (T-Per1-19957; eclipsing binary) 2MASS J08511335+1151401 (HU Cnc; RS Canum Venaticorum-type system), 2MASS J08511799+1145541 (HV Cnc; eclipsing binary) 2MASS J08512079+1153261 (ES Cnc; eclipsing binary), 2MASS J08512814+1149274 (EV Cnc; eclipsing binary), 2MASS J20313481-4916270 (ASAS J203135-4916.5, Delta-Scuti type variable) 2MASS J22004033+2056425 (GSC 01692-01074; eclipsing binary, found by Patrick Wils using the 2MASS Cal-PSWDB).

In the ρ Ophiucus molecular cloud region 21 stars were found in our search for dimming events. This region has been comprehensively studied previously for variability with the same data (Parks et al. 2013), and in this region, we only found dimming events in highly variable stars that were a subgroup of those previously identified by Parks et al. (2013). We found no quiescent star in the region, spending most of its time with little or no variability, that exhibited a dimming event. We defer the discussion of these objects to section 5.

3.1 Periodic stars

For stars with candidate dimming events that were not previously identified as variable stars, (and excluding the ρ -Ophiucus region) we searched for periodicity. We computed the frequency spectrum of the light curve and then examined the phased or period-folded light curves at the locations of peaks or multiples of peaks in the spectrum. Periods were measured with the software package Peranso Light Curve and Period Analysis Software⁶ or with the Period04 Java package⁷. We checked for internal consistency and estimated errors using both software packages.

46 new eclipsing binary candidates and 6 stars with periodic light curves (intrinsic pulsators such as RR-Lyrae type variables) were found in this search. Phased or period-folded light curves for these stars are shown in Figure 1 (eclipsing binaries) and Figure 2 (intrinsic pulsators). Periods for these stars are listed with their 2MASS J identifier in Tables 1 and 2. In these tables, J -band photometry and $J - K$ color were taken from the 2MASS Point Source Catalog. Estimated primary eclipse depth, in J magnitudes, and the Julian dates of a primary eclipse are listed in 1 for the eclipsing binaries. Tentative eclipsing binary types EA (Algol type; detached) or EB (β Lyrae type; semi-detached) or EW (W Ursae Majoris type or contact) were assigned based on the shape of the phased light curves. For the intrinsic pulsators, peak to peak amplitudes in J band and the Julian date of a minimum are listed in Table 2. We also list colors computed from mid

³ The VizieR database is hosted and developed by Centre de Données astronomiques de Strasbourg <http://vizier.u-strasbg.fr/viz-bin/VizieR>

⁴ Watson, C., Henden, A.A., & Price, A. 2006–2014 <http://www.aavso.org/vsx>

⁵ <http://www.astronet.ru/db/varstars/msg/1234323>

⁶ <http://www.peranso.com/> by Tonny Vanmunster of the CBA Belgium Observatory

⁷ P. Lenz and M. Breger, 2005, Period04 User Guide, Communications in Asteroseismology, 146, 53

infrared photometric measurements from the Wide-Field Infrared Survey Explorer (WISE) (Cutri et al. 2012; Wright et al. 2012). The WISE observatory mapped the sky at 3.4, 4.6, 12 and 22 μm and these are referred to as the W1, W2, W3 and W4 bands, respectively. Zero points for the bands in Jy are given by Jarrett et al. (2011).

From Table 1 a couple of the eclipsing binary candidates (e.g., 2MASS J03314787+3746510) have red enough colors that they could have been in the sample searched by Plavchan et al. (2008), however they were excluded because of their low galactic latitude (Plavchan et al. 2008 excluded $|b| < 20^\circ$). Because of their low extragalactic latitude, it is not obvious whether any of the red eclipsing binaries are compared of late type stars (e.g., Becker et al. 2008). A few of the eclipsing binaries, 2MASS J03314787+3746510, 2MASS J12014395-4958101, 2MASS J18391249+4845394 are contact eclipsing binaries with short periods $P \sim 0.25$ days. They are contact main sequence eclipsing binaries or W UMa systems. They are of interest as distance indicators and for constraining stellar models (see for example Dimitrov & Kjurkchieva 2010; Norton et al. 2011; van Eyken et al. 2011; Nefs et al. 2012). The periods of these three contact eclipsing binaries are not so short as to make them exceptional compared to those previously discovered. 2MASS J16265638+0552039 has a period of 0.339974 days but it is not a contact binary or a W Uma system, it is detached. Its galactic longitude is 34 degrees implying negligible extinction in the near-IR bands and its color fairly red, $J - K = 0.92$ mag. The infrared colors would be consistent with an M3 and not later spectral type (placing it on Figure 11 by van Eyken et al. 2011). This system is similar to the interesting short period eclipsing M-dwarf binaries, including detached binary systems, recently found by Nefs et al. (2012).

We looked in the EB sample to see if any had interesting infrared excesses and so were likely to be young. There were two with WISE mid-infrared fluxes, 2MASS J19015103-0419487 and 2MASS J19020201-0430160 listed in the all sky catalog (Cutri et al. 2012). However both are near the Galactic plane and inspection of the WISE imaging atlas implied that there was confusion with mid-infrared extended emission.

Four of the periodic variables that do not resemble eclipsing binaries have short periods (less than 4 days) and folded-light curves resembling RR-Lyrae type variables. However, a semi-regular or long period variable star was found, 2MASS J19015057-0431087, with a period of ~ 220 days. This object has red WISE colors $[3.4]-[4.6]=0.41$ (Vega magnitudes) and $[4.6]-[12]=0.77$ mag, implying that the object has a dusty envelope. Its period and mid-infrared colors are similar to the semi-regular or long period variables discussed by Palaversa et al. (2013) that could be AGB stars. The scatter in the period-folded light curve is larger than the photometric errors because the light curve is not reproducible between different cycles. This behavior was also seen in the long period variables identified by Palaversa et al. (2013) (see their figure 16).

3.2 Additional variable objects

Some objects exhibited deep dimming events, but we could not identify periods for them, either because they exhib-

ited variability on many timescales or we lacked sufficient number of eclipse-like features. Excluding the ρ Ophiucus region, these objects are listed in Table 3 and their light curves are shown in Figure 3 or if they are extragalactic in Figure 4. 2MASS J12141394+3541567 and 2MASS J16311840+2946322 are extragalactic objects based on their membership in the galaxy catalog by Paturel et al. (2003). On the left hand side in these figures, the day averages are plotted as black points on top of individual measurements (colored points). From top to bottom are the $J - K$ color, J , H , and K magnitudes as a function of time. On the right hand side we show color-magnitude plots (J vs $J - K$) but with each point corresponding to a different observation time. On the color-magnitude plots, we also show a reddening vector with length corresponding to an extinction of $A_J = 0.1$ (using the standard extinction law in Table 1 by Mathis 1990).

Most of the non-extragalactic aperiodic objects exhibiting deep dimming events have WISE colors consistent with stellar sources. The object 2MASS J17482764-4531581 has W2-W3 color $[4.6\mu\text{m}] - [12\mu\text{m}] = 2.51$ (Vega magnitudes), and 2MASS J08254738-3906306 has a color $[4.6]-[12]$ of 1.75 mag. These red colors are more typical of a Seyfert galaxy than a star (placing it on Figure 14 by Yan et al. 2013) but would also be typical of a Class I young stellar object (see Figure 2 by Wolk et al. 2013). Using the WISE image server, we found that the $12\mu\text{m}$ flux for 2MASS J17482764-4531581 is confused with a nearby shell, so the red $[4.6] - [12]$ color is probably not that of the object. 2MASS J19020170-0405561 with $[4.6]-[12]\mu\text{m} = 0.80$ mag also has a flux at $12\mu\text{m}$ and $22\mu\text{m}$ of 8.3 and 10 mJy, respectively. We WISE images showed that there was a source at 12 and 22 μm . This object likely has an infrared excess at $22\mu\text{m}$ even though its $[4.6]-[12]\mu\text{m}$ color is only moderately red. Based on its mid-infrared excess, this object could be an active galaxy or a young stellar object.

Even though 2MASS J19020170-0405561 and 2MASS J19020170-0405561 are not identified as extragalactic, their brightness and color variations are strongly correlated, similar to the two galaxies, 2MASS J12141394+3541567 and 2MASS J16311840+2946322. 2MASS J18394250+4856177 also exhibits correlated brightness and color variations but has no sign of infrared excess. It is common for young stellar objects to exhibit transient dimming events (Morales-Calderon et al. 2011; Cody et al. 2014; Parks et al. 2013). 20% of the young disk-bearing stars in NGC2264 surveyed by Cody et al. (2014) were classified as “dippers”. Dimming in young stellar objects is often, but not always, associated with reddening and so due to extinction (Cody et al. 2014). 2MASS J08254738-3906306, and 2MASS J19020170-0405561, with infrared excesses, might be young stellar objects rather than active galaxies. 2MASS J18394250+4856177, exhibiting correlated color/magnitude variations but no infrared excess, is difficult to classify, even though it is not faint, at 12.7 mag in J band.

2MASS J19014896-0453423 is a bright object (9.8 mag at J band) varying by 0.25 magnitudes, but it does not display strong color variations. It clearly has an infrared excess as it was detected with WISE at $22\mu\text{m}$ with a flux of 11 mJy. Unfortunately, it was not observed with WISE at $12\mu\text{m}$, making it difficult to classify based on its mid-infrared color.

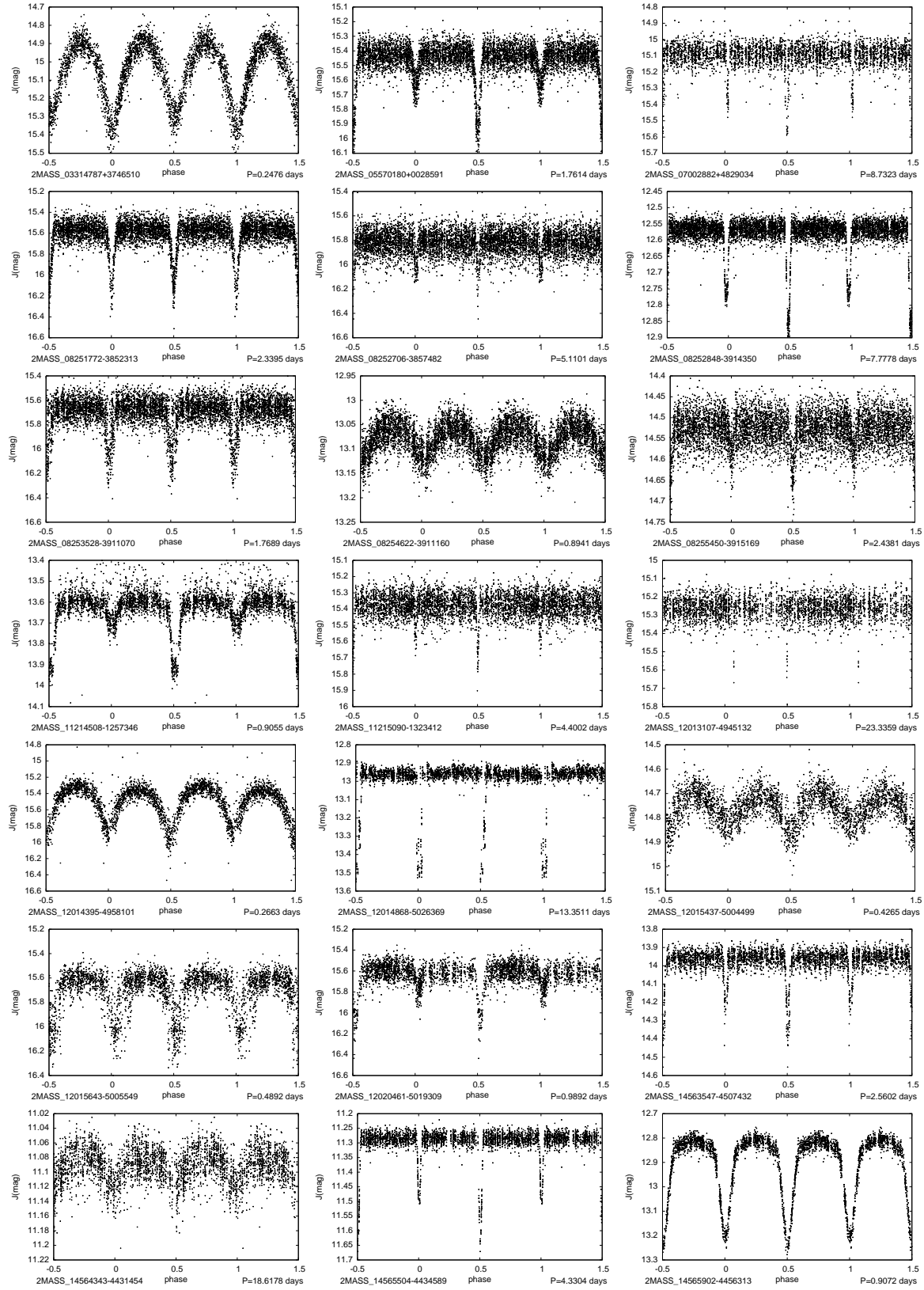


Figure 1. Phased (period-folded) light curves of eclipsing binary candidates in J-band found in a search for dimming events.

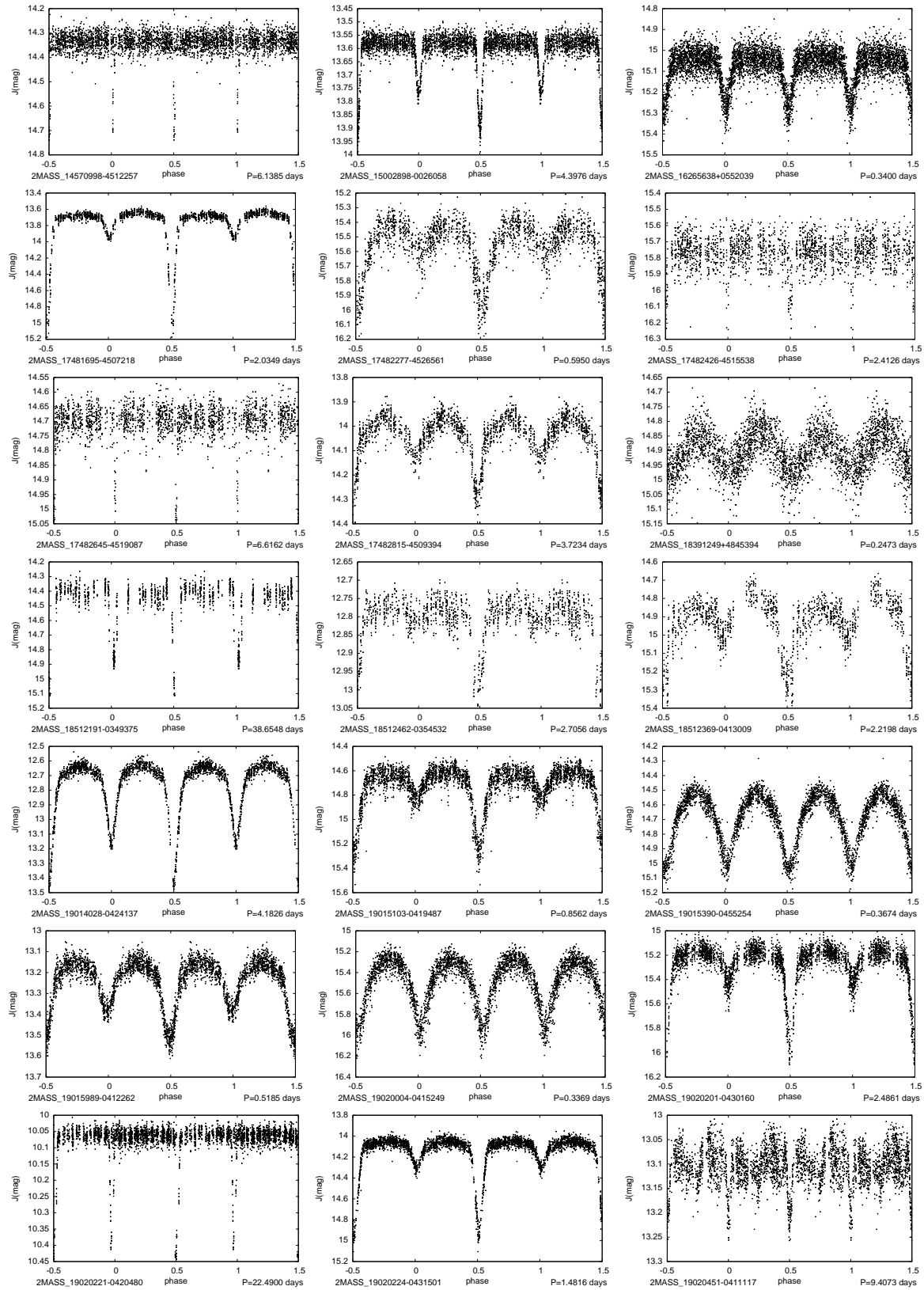


Figure 1. Figure continued.

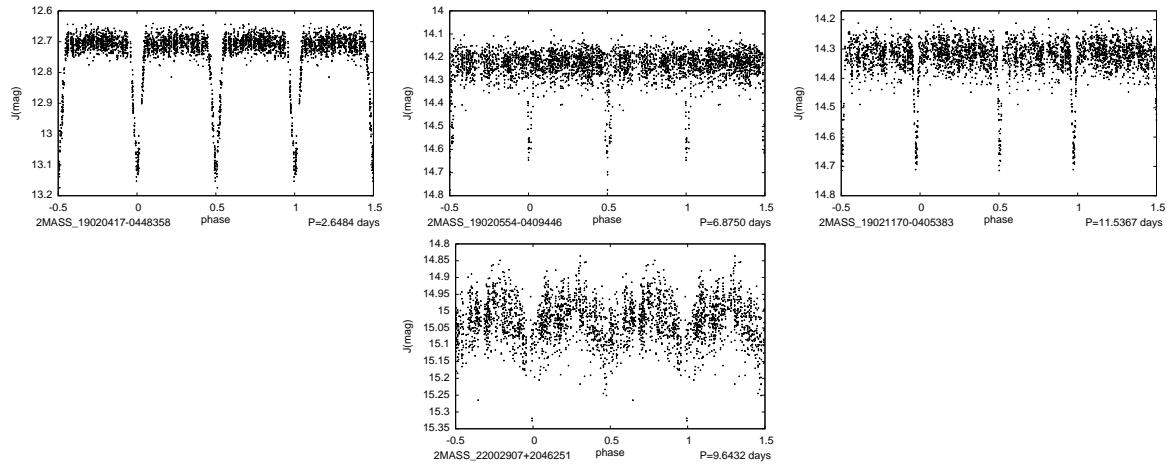


Figure 1. Figure continued.

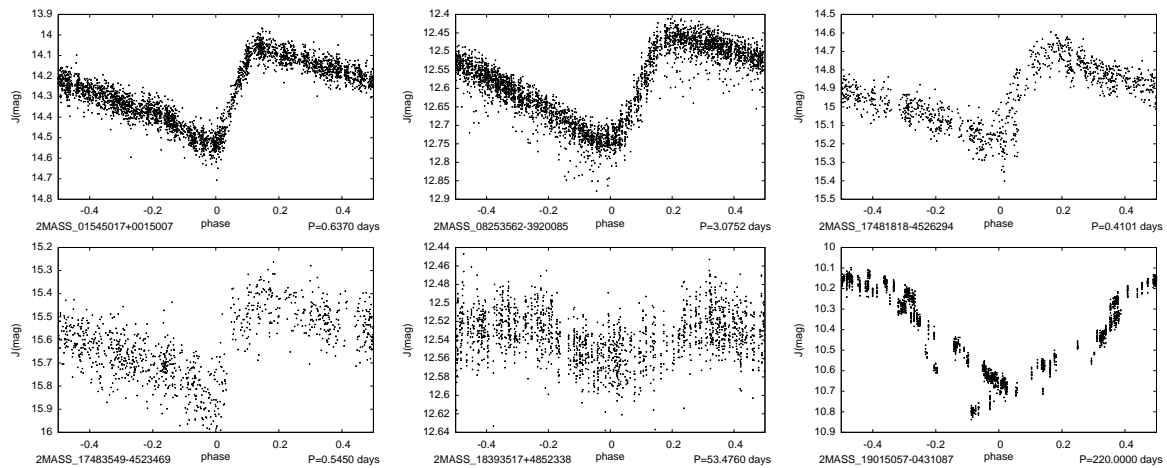


Figure 2. Phased (period-folded) light curves (similar to Figure 1) in J -band of previously unknown periodic stars (excluding eclipsing binaries) found in a search for dimming events. 2MASS J01545017+0015007, 2MASS J08253562-3920085, 2MASS J17481818-4526294 and 2MASS J17483549-4523469 are probably RR-Lyrae type variables. 2MASS J18393517+4852338 could be a Cepheid type variable and 2MASS J19015057-0431087 is a long period variable with an infrared excess that is not precisely periodic.

The remaining objects, 2MASS J08321728-0114337, 2MASS J11213899-1250040, 2MASS J17480955-4459033, 2MASS J17482764-4531581, and 2MASS J22001106+2032472, display isolated dimming events that do not strongly change the color. The lack of color variation during dimming events implies that extinction (from interstellar dust) is not the cause of the dimming. We examined the data points during each dimming event but did not detect sub-structure in the shape of the light curve. They are either dimming, as if at the beginning of an eclipse, flat, or brightening, as if at the end of an eclipse. In most cases we don't actually know the length of the eclipse. If these dimming events are real, then these objects are likely to be stellar eclipsing binaries, and we have not been able to identify their periods because we have not seen enough eclipses. An object completely occulting a star on a day-long timescale must move a speed of order 10km/s (assuming approximately a solar radius for the primary). Only one of these, 2MASS J17480955-4459033 has an extremely deep dimming event, at 0.6 mag, implying

that a significant fraction of the starlight is obscured. The occulter could be molecular cloud material (Rodriguez et al. 2013) rather than a secondary star or a disk.

To summarize, our color insensitive but sigma-limited search for dimming events primarily found eclipsing binaries. Including the previously known eclipsing binaries we found a total of 59 eclipsing binaries from approximately 40000 stars implying that approximately 1/700 stars searched is an eclipsing binary. We also found 6 periodic or quasi periodic intrinsic pulsators. An 11 additional objects were found with aperiodic dimming events. Of these two are extragalactic objects and so probably active galaxies. A few exhibited dimming events without strong color variation. These could be eclipsing binaries but we have failed to see enough eclipses to identify a period. A few exhibit correlations between dimming and color suggesting that they are active galaxies or young stellar objects. The 11 objects found exhibiting dimming events could be interesting objects, but further observations are needed to classify them. Excluding

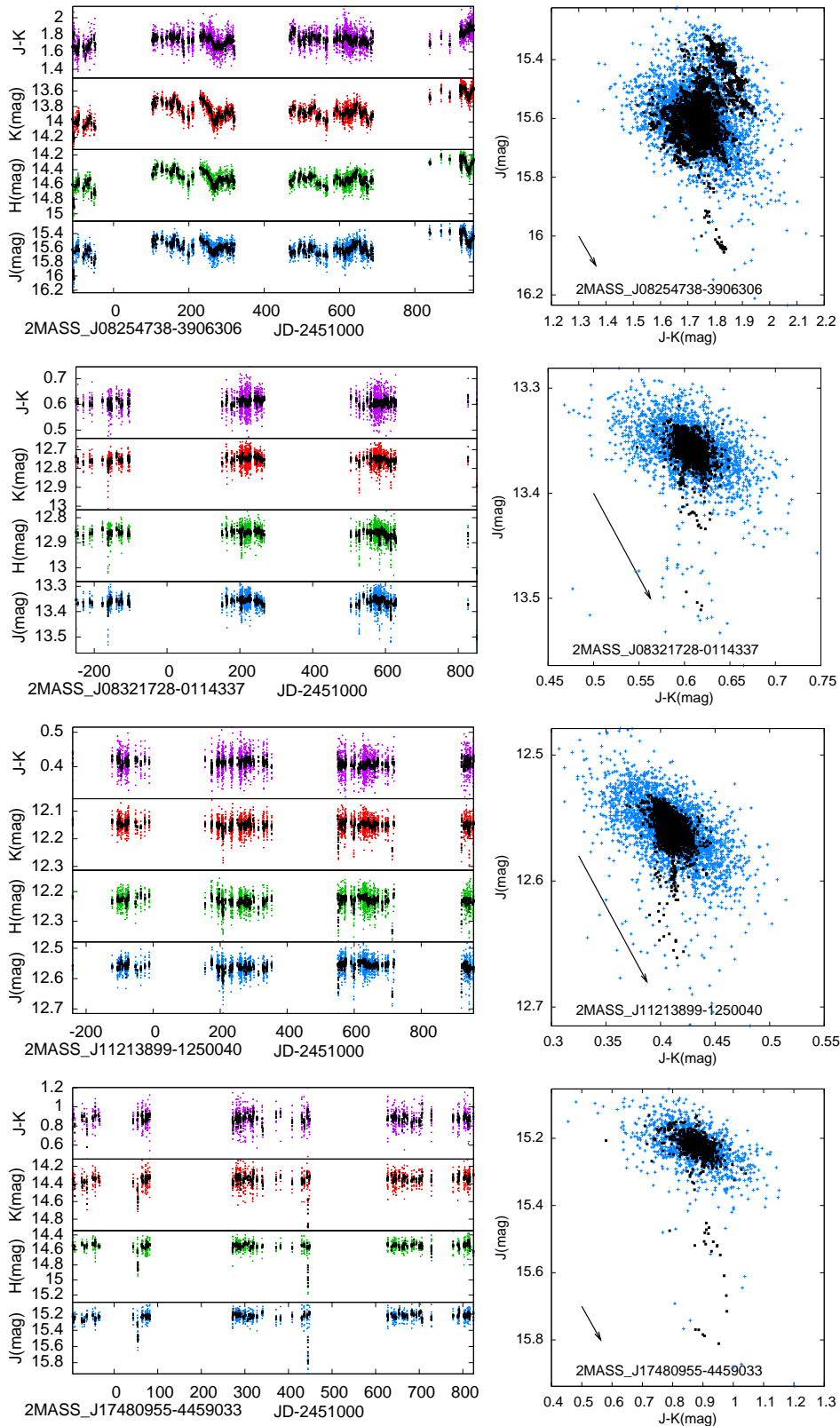


Figure 3. Variable objects listed in Table 3 with deep dimming events. We failed to find periods for these objects. The entire light curves are shown in J, H, K and in $J - K$ magnitudes as a function on time on the left. On the right we show a color magnitude plot with J vs $J - K$ magnitudes at different times. A reddening vector with length $A_J = 0.1$ is shown on the lower left in the color-magnitude plots. Photometric measurements are shown as colored points. Box-day averages are shown as black points. Some of these objects do display significant color variations during dimming events, suggesting that these are eclipsing binary stars but we have not observed enough eclipses to identify a period. For objects with correlated color/magnitude variations, they could be active galaxies or young stellar objects.

Table 1. Eclipsing binary candidates

Object	J (mag)	$J - K$ (mag)	Period (days)	Depth (mag)	t_{min} (days)	Type
(1)	(2)	(3)	(4)	(5)	(6)	(7)
2MASS J03314787+3746510	15.248±0.056	1.01	0.2475799±9e-06	0.5	0.204	EW
2MASS J05570180+0028591	15.454±0.051	0.71	1.761380±0.000418	0.7	2.06	EA
2MASS J07002882+4829034	15.155±0.044	0.51	8.732304±0.014124	0.4	2.52	EA
2MASS J08251772-3852313	15.565±0.067	0.56	2.33949±0.00014	0.8	1.27	EA
2MASS J08252706-3857482	15.824±0.042	0.80	5.1101 ±0.0005	0.6	2.70	EA
2MASS J08252848-3914350	12.550±0.026	0.36	7.777814±0.0068	0.4	3.9	EA
2MASS J08253528-3911070	15.565±0.066	0.28	1.76891±0.00011	0.5	1.35	EA
2MASS J08254622-3911160	13.149±0.028	0.92	0.89405± 0.00016	0.1	0.77	EW
2MASS J08255450-3915169	14.534±0.032	1.39	2.43813±0.00030	0.15	2.77	EA
2MASS J11214508-1257346	13.597±0.030	0.46	0.90553±0.0001	0.35	0.60	EB
2MASS J11215090-1323412	15.373±0.042	0.82	4.40021±0.0016	0.5	2.26	EA
2MASS J12013107-4945132	15.317±0.052	0.32	23.3359±0.0049	0.4	14.67	EA
2MASS J12014395-4958101	15.421±0.056	0.8	0.266297±7e-06	0.7	0.08	EW
2MASS J12014868-5026369	12.957±0.027	0.68	13.3511±0.00890	0.4	3.2	EA
2MASS J12015437-5004499	14.729±0.051	0.52	0.426512±5e-05	0.55	0.45	EA
2MASS J12015643-5005549	15.515±0.063	0.76	0.489237±2e-05	0.5	0.50	EB
2MASS J12020461-5019309	15.645±0.060	0.65	0.989157±0.000053	0.2	0.45	EB
2MASS J14563547-4507432	13.955±0.032	0.84	2.560226± 0.00034	0.45	2.18	EA
2MASS J14564343-4431454	11.094±0.021	0.88	18.618±0.041	0.1	9.31	EW
2MASS J14565504-4434589	11.294±0.022	0.32	4.330373±0.001351	0.35	4.43	EA
2MASS J14565902-4456313	12.824±0.026	0.20	0.907237±0.000006	0.4	1.14	EB
2MASS J14570998-4512257	14.508±0.032	0.51	6.1384±0.0120	0.35	2.47	EA
2MASS J15002898-0026058	13.629±0.032	0.50	4.397585±0.001950	0.35	4.40	EA
2MASS J16265638+0552039	15.065±0.050	0.92	0.339974 ± 0.000004	0.3	0.148	EA
2MASS J17481695-4507218	13.713±0.029	0.50	2.03489±0.0003	1.3	2.17	EB
2MASS J17482277-4526561	15.438±0.048	0.46	0.595025±7e-05	0.8	0.357	EB
2MASS J17482426-4515538	15.724±0.067	0.82	2.412559±0.00050	0.4	0.66	EA
2MASS J17482645-4519087	14.732±0.032	0.44	6.61616±0.00069	0.35	8.24	EA
2MASS J17482815-4509394	14.064±0.029	0.65	3.72342±0.00097	0.4	3.06	EB
2MASS J18391249+4845394	14.815±0.037	0.51	0.2473227±2e-05	0.15	0.140	EW
2MASS J18512191-0349375	14.362±0.028	1.39	38.6548±0.116	0.6	31.6	EA
2MASS J18512369-0413009	14.812±0.029	1.82	2.21975±0.00049	0.5	2.08	EB
2MASS J18512462-0354532	12.828±0.045	0.58	2.70563 ±0.00073	0.25	2.95	EB
2MASS J19014028-0424137	12.687±0.032	0.95	4.182611±0.0003	0.8	1.89	EB
2MASS J19015103-0419487	14.661±0.026	0.61	0.856174±1e-05	0.7	0.13	EB
2MASS J19015390-0455254	14.893±0.041	0.62	0.36737±7e-06	0.6	0.26	EW
2MASS J19015989-0412262	13.175±0.029	0.64	0.518462±1e-05	0.4	0.46	EB
2MASS J19020004-0415249	15.458±0.084	0.66	0.336916±2e-06	0.9	0.22	EW
2MASS J19020201-0430160	15.322±0.067	0.97	2.486081± 0.00035	0.9	0.84	EB
2MASS J19020221-0420480	10.074±0.026	0.29	22.49±0.36	0.35	19.85	EA
2MASS J19020224-0431501	14.096±0.036	0.64	1.481571±9e-05	0.9	1.02	EB
2MASS J19020417-0448358	12.703±0.026	0.48	2.64844±0.00014	0.45	3.14	EA
2MASS J19020451-0411117	13.148±0.024	0.87	9.4073±0.0026	0.15	6.90	EA
2MASS J19020554-0409446	14.271±0.030	0.59	6.87500±0.00347	0.5	3.3	EA
2MASS J19021170-0405383	14.276±0.037	0.47	11.5367±0.004	0.4	6.17	EA
2MASS J22002907+2046251	15.041±0.052	0.69	9.643±0.019	0.15	9.64	EW

Here we list previously unknown eclipsing binaries found in a search for dimming events. Columns: (1) Identifier from the 2MASS point source catalog. (2,3) J magnitudes and colors are from the 2MASS point source catalog (Cutri et al. 2003). (4) The estimated period in days. (5) The estimated depth of primary eclipse is given here in J magnitudes. (6) Epoch of primary eclipse. We list the Julian Date in days subtracted by 2450000.0. (7) Binaries are roughly classified as detached (EA), experiencing gravitational deformation (EB), or in contact (EW), based on the shapes of the light curves.

the young stellar objects, we failed to find an object that dimmed longer than a single night.

4 A SEARCH FOR RED DIMMING EVENTS

Eclipsing binaries usually do not exhibit strong color variations during eclipses. However, transient extinction would change the color of a star during a dimming event. For example, OGLE-LMC-ECL-11893 disk eclipses are 0.3 magnitude deeper in B band than I band (Dong et al. 2014; see their

Table 2. Periodic stars that are not eclipsing binaries

Object (1)	J (2)	$J - K$ (3)	Period (4)	Amplitude (5)	t_{min} (6)	[3.4]-[4.6] (7)	[4.6]-[12] (8)
2MASS J01545017+0015007	14.626±0.041	0.48	0.636980±2e-05	0.3	0.318	0.00	<1.49
2MASS J08253562-3920085	12.568±0.030	1.36	3.07517 ±0.00081	0.5	0.938	0.04	0.58
2MASS J17481818-4526294	14.806±0.033	0.12	0.4101433±2e-05	0.5	0.285	-0.90	<3.2
2MASS J17483549-4523469	15.353±0.060	0.50	0.545027± 3e-05	0.5	0.5225	-	-
2MASS J18393517+4852338	12.537±0.021	0.85	53.476±0.87	0.08	41.74	-0.06	0.043
2MASS J19015057-0431087	10.382±0.205	1.82	220±30	0.55	0.0	0.41	0.77

Here we list previously unknown periodic variable stars (excluding eclipsing binaries) found in a color insensitive search for dimming events. Columns are similar to those of Table 1 but listing variable stars likely to be intrinsic pulsators. The period folded light curves are shown in Figure 2. Columns: (5) The amplitude (peak to peak) of variation in J band magnitude. (6) Epoch of a minimum in the light curve. We list the Julian Date in days subtracted by 2450000.0. (7, 8) Colors in Vega magnitudes computed from broad band photometry from the WISE satellite (Cutri et al. 2012; Wright et al. 2012). For 2MASS J01545017+0015007 and 2MASS J17481818-4526294 the [4.6]-[12] colors are upper limits as the sources were not detected at 12 μm . 2MASS J17483549-4523469 was not in the WISE all sky catalog.

Table 3. Variable object candidates without measured periods and with dimming events

Object (1)	J (2)	σ_J (3)	H (4)	σ_H (5)	K (6)	σ_K (7)	$J - K$ (8)	[3.4]-[4.6] (9)	[4.6]-[12] (10)	R (11)	b (12)
2MASS J08254738-3906306	15.586	0.116	14.493	0.128	13.833	0.142	1.75	0.49	1.75	0.20	0.221±0.019
2MASS J08321728-0114337	13.360	0.028	12.863	0.030	12.753	0.036	0.61	-0.05	0.13	0.38	0.304±0.016
2MASS J11213899-1250040	12.560	0.025	12.231	0.027	12.150	0.027	0.41	-0.03	0.54	0.52	0.459±0.015
2MASS J17480955-4459033	15.228	0.074	14.548	0.075	14.350	0.091	0.88	-0.18	< 2.5 ^a	0.44	0.350±0.023
2MASS J17482764-4531581	12.929	0.076	12.599	0.077	12.524	0.077	0.41	-0.06	2.5 ^b	0.12	0.285±0.074
2MASS J18394250+4856177	12.623	0.073	12.142	0.090	12.080	0.057	0.54	-0.09	0.05	0.62	0.830±0.025
2MASS J19014896-0453423	9.786	0.059	8.780	0.053	8.440	0.049	1.35	0.05	0.5 ^c	0.56	1.189±0.046
2MASS J19020170-0405561	11.453	0.044	10.339	0.044	9.913	0.038	1.54	0.02	0.8	0.53	0.929±0.035
2MASS J22001106+2032472	13.069	0.025	12.735	0.032	12.665	0.033	0.40	-0.04	< 0.2 ^a	0.42	0.321±0.029
2MASS J12141394+3541567 ^d	16.151	0.140	15.456	0.154	14.902	0.145	1.25	0.28	3.64	0.59	0.464±0.012
2MASS J16311840+2946322 ^d	15.185	0.110	14.526	0.135	14.057	0.115	1.13	0.29	3.97	0.43	0.436±0.027

Here we list objects found in a color insensitive search for dimming events that lack measured periods. In this table we have excluded young stellar objects in the ρ Ophiucus region. Objects are sorted into two groups, with the second group identified as extragalactic objects. Columns: (1): 2MASS point source catalog identifier. (2)-(7): Mean and standard deviations, in magnitudes, computed from the light curves in J , H , and K bands, respectively. (8): $J - K$ Color in magnitudes was computed from the mean magnitudes of the light curve. (9,10): Colors in Vega magnitudes from the WISE survey (Cutri et al. 2012). The four bands of the WISE survey bands are centered at 3.4, 4.6, 12 and 22 μm . (11): The cross correlation coefficient, R , (equation 1) between color and magnitude. (12): The slope, b , of the best fitting line on the color/magnitude plot (equation 2).

^a For 2MASS J17480955-4459033 and 2MASS J22001106+2032472 the [4.6]-[12] color is an upper limit as the sources were not detected at 12 μm .

^bThe WISE image server showed that the 12 μm flux for 2MASS J17482764-4531581 is confused with a nearby shell, so the red [4.6 μm] - [12 μm] is probably not that of the source.

^c2MASS J19014896-0453423 was not observed at 12 μm with the WISE satellite, however the source was detected at 22 μm . The [4.6]-[12] color is estimated from the fluxes at 4.6 and 22 μm .

^d2MASS J12141394+3541567 and 2MASS J16311840+2946322 are extragalactic objects based on their membership in the galaxy catalog by Paturel et al. (2003).

Figure 3) and EE-Cep changes color during eclipse (Galan et al. 2012). Even though OGLE-LMC-ECL-11893 was first identified as an eclipsing binary, its light curve during eclipse is complex (asymmetric and W-shaped) and so is likely due to a disk that is hosted by the secondary star. For a standard Milky Way dust extinction law (Mathis 1990, Table 1) the change in color $E(J-K) \sim 0.625A_J$, with A_J the extinction in J band and both A_J and $E(J-K)$ in magnitudes. During a dimming event caused by extinction from interstellar-like dust, we expect an increase in $J - K$ color about half of the depth in magnitudes of the event in J band.

We carried out a search independent of the magnitude dispersions (of the entire light curves), by searching for dimming events in J band greater than 0.15 mag that are ac-

companied by reddening in color $J - K$ of greater than 0.07 with both values compared to the means, μ_J , and $\mu_J - \mu_K$, of the entire light curve. These levels were chosen to yield a moderate (and not extremely large) number of objects, exhibiting dimming, with statistically significant color variations during dimming. Because we are looking for small color changes, we used the day averaged light curves for this search. We discarded lower quality data points with a sum of standard deviations $\sigma_{Jb} + \sigma_{Kb} > 0.08$, where σ_{Jb}, σ_{Kb} are the uncertainty of the box-averaged light curve in J and K bands, respectively. We discarded events with nearby source confusion, and sparse data (a particular problem for the 90547 tile) and at times where more than 1 star in the same calibration tile exhibited an event (and implying that

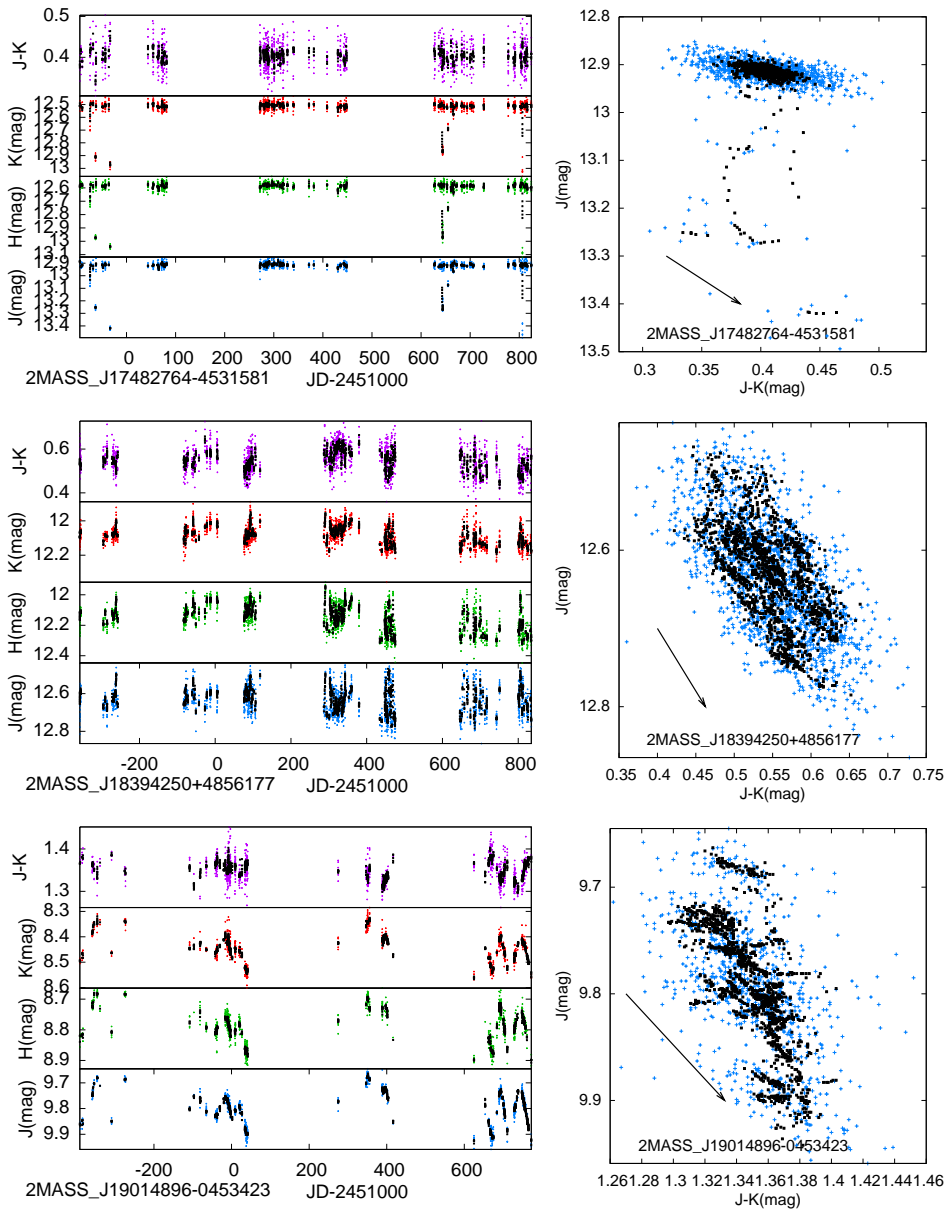


Figure 3. Figure continued.

the dimming was due to weather rather than intrinsic to the sources themselves). Our search criteria do not depend on the standard deviation of the entire light curve so we are sensitive to strongly variable objects.

4.1 Active Galaxies

This search revealed a number of variable objects in which the object systematically is bluer when brighter. The objects that are extragalactic objects, we have listed in Table 4. Table 4 also lists mean magnitudes and standard deviations of the magnitude distributions we measured from the entire light curves. An object was designated extragalactic if we found it in a galaxy catalog such as Paturel et al. (2003, 2005) or it was classified as a galaxy in the Sloan Digital Sky Survey (SDSS) or found in the 2MASS extended source

catalog (see the references listed in column 13 and notes in Table 4). An additional object we classify as extragalactic based on its mid-infrared color as measured with WISE. This object, 2MASS J08255140-3907590, is quite red, with $J-K \sim 2.2$ and nearly in the Galactic plane, with a Galactic latitude of only -0.65 degrees, but it is not resolved, found to be as extended nor listed in any extragalactic catalog. However, its mid-infrared flux is 0.028 Jy at $12 \mu\text{m}$ and 0.06 Jy at $22 \mu\text{m}$ (based on the WISE measurements) and is increasing with increasing wavelength between 12 and 22 microns. Mira variables and young stellar objects that can be seen in the near-infrared tend to have flux decreasing with increasing wavelength in the $12-22 \mu\text{m}$ region. The 3.4 and $4.6 \mu\text{m}$ color, based on the WISE W1 and W2 bands, respectively, $[3.4] - [4.6] = 1.05$ (in Vega magnitudes), is consistent with $[3.4] - [4.6] > 0.7$ mag found for quasars (Yan et al. 2013). Consequently 2MASS J08255140-3907590

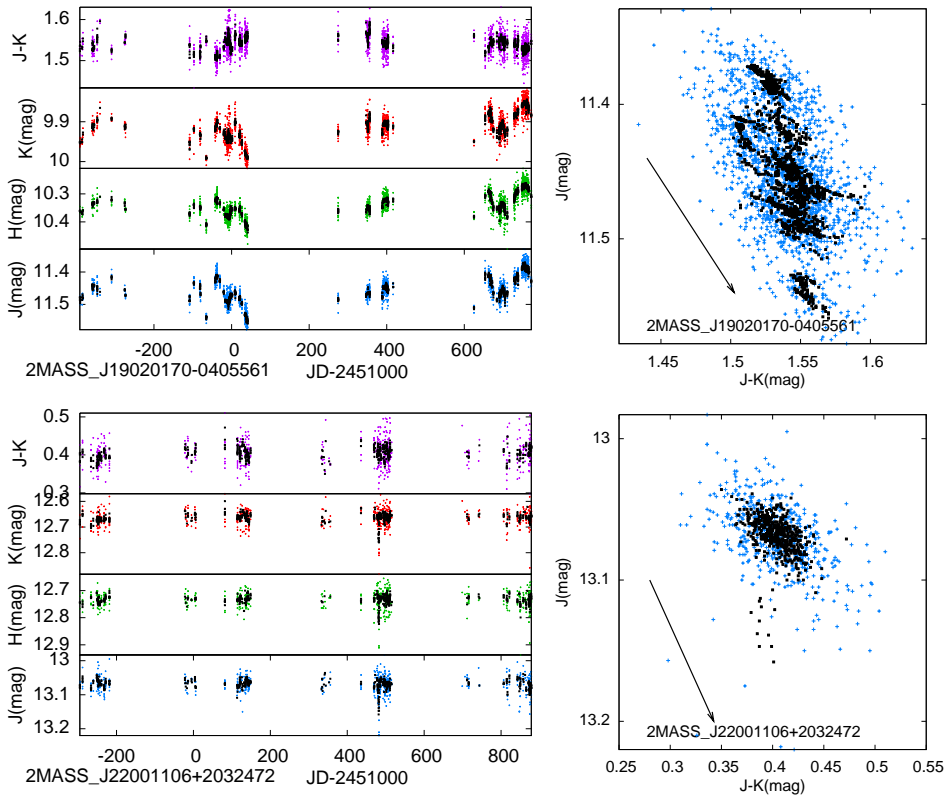


Figure 3. Figure continued.

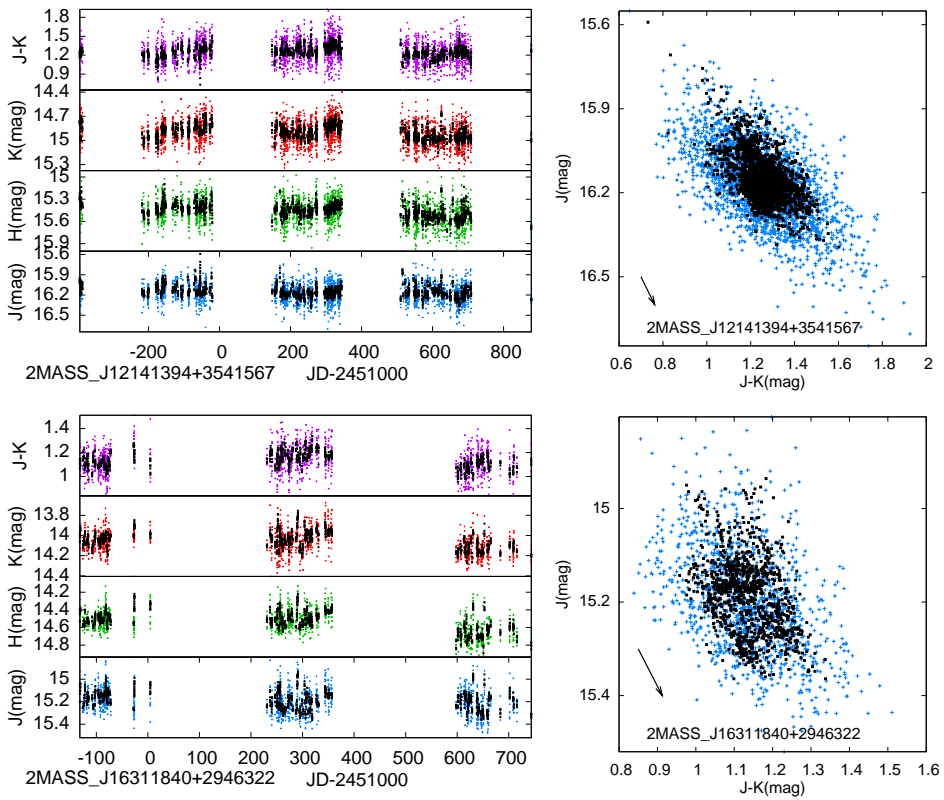


Figure 4. Similar to Figure 3 but for extragalactic variable objects listed in Table 3 with deep dimming events.

is most likely an active galaxy and we have designated it as extragalactic.

Fainter AGNs vary more strongly than their bright counterparts and they vary more strongly on shorter wavelengths (e.g., di Clemente et al. 1996). Brightness and color are correlated with objects becoming bluer when brighter. This is usually attributed to a contribution to the flux in red light from non-variable emission source (narrow line region, galaxy or star light) combined with a variable bluer component associated with the active nucleus (Winkler 1997; Webb & Malkan 2000; Sakata et al. 2010). Since active galaxies become bluer when they are brighter, they also are redder when they are fainter. This explains why we mostly found faint AGNs when searching for red dimming events. The object, 2MASS J08255140-3907590 (that is bright in the mid-infrared) is variable on longer timescales than the other extragalactic objects, suggesting that it is a more luminous active galaxy than the others, perhaps a quasar.

The color magnitude plots for the active galaxies (shown in Figure 5) show that the the color and magnitude variations are strongly correlated. However the distribution in these plots is often bowed, with less color variation taking place when the object is brightest and more color variation when the object is faintest (e.g., the color magnitude plot of 2MASS J0341750+0646008). In some cases, color magnitude variations are not restricted to a single one-dimensional line (e.g., 2MASS J00241907-0138176) implying that there is more than one process contributing to the color and brightness variations. For example, there might be both flares and extinction.

4.2 Additional Objects

An additional two objects were found in the red dimming search for which we don't have classification. These are shown in Figure 6 and listed in Table 5. 2MASS J18395818+4903590 was not observed by WISE so we have no information on its mid-infrared colors. 2MASS J07004330+4828115 has a red [4.6]-[12] color of 3.03 mag suggesting that the object is extragalactic.

Three eclipsing binaries were found in the red-dimming search that were also found in the color insensitive search. We discuss these objects below in section 6.1 where we discuss the objects found in both dimming searches.

To summarize, after excluding the ρ Ophiucus region, the red dimming search found three eclipsing binaries (discussed below), 14 active galaxies, and 3 additional objects that exhibit correlations between color and brightness similar to those seen in the active galaxies (one of these was also found in the color insensitive search and was previously discussed).

5 YOUNG STELLAR OBJECTS FROM THE ρ OPHIUCUS REGION

Our dimming searches revealed 21 objects in the ρ Ophiucus molecular cloud region that were previously studied by Parks et al. (2013). These young stellar objects are listed in Table 6 separated into two groups, 11 found only in the color insensitive search, followed by 10 found in both searches. Their color magnitude distributions are shown in

Figures 7 and 8. These young stellar objects are comparable to stars exhibiting brief, sharp drops in stellar flux, denoted as AA Tau (Herbst et al. 1994) or dipper variables, seen in the Orion Nebula and NGC 2264 stellar clusters (Morales-Calderon et al. 2011; Cody et al. 2014) (also see light curves from the Cygnus OB7 association Wolk et al. 2013).

Among the young stellar objects found in the red-dimming search, all but 1 of the 10 young stellar objects found in the red-dimming search were found to be periodic by Parks et al. (2013) and all but 1 of the 10 are Class II objects, but since we used both J and K band for the search we necessarily excluded extremely red young stellar objects. The exception is YLW 16b, classified as a long timescale variable and a Class I young stellar object (Parks et al. 2013). Here young stellar object (YSO) classes were assigned based on those by Bontemps et al. (2001); Gutermuth et al. (2009). Among the young stellar objects only found in the color insensitive dimming search, all are Class II objects except ISO-Oph 135 which is a Class III object. The objects found only in the color insensitive search included irregular and long timescale variables as well as periodic variables.

Parks et al. (2013) classified the periodic behavior based on the color variations seen in the period-folded light curves. They classified the periodic behavior as eclipse-like, inverse eclipse like, or sinusoidal with hot or cold star spots. If dimming was associated with redder colors, then extinction was assumed, whereas if dimming was not associated with redder colors, periodicity was associated with cool or hot star spots, depending upon the fraction of phase spent bright compared to dim. The aperiodic variables were divided into long time scale variables (LTVs) and irregulars and their brightness variations classified as due to accretion if they were colorless or extinction if dimming was associated with reddening. These classifications are listed in our Table 6. We expected to only find eclipse-like or extinction-like classifications in our red dimming event selected sample, however some of them were classified as having inverse-eclipse-like period-folded light curves and some with sinusoidal with colors variations expected from star spots. Our color magnitude diagrams show that *all* the young stellar objects predominantly vary along the reddening direction, even those only found in the color insensitive search. The objects with the largest variations in both color and brightness tend to be those that have the strongest correlations between reddening and brightness and these were found in the red dimming search. The color magnitude distributions are not bowed as were some of the color/magnitude distributions of the active galaxies. Not all the young stellar objects displayed strong correlations in color vs magnitude, for example Iso-Oph 151 (2MASS J16273084-2424560) has a broad color magnitude distribution. While features are seen in the distribution along the reddening direction the distribution is broad implying that more than one emission process is affecting the brightness and color.

One of the young stellar objects, 2MASS J16271848-2429059, WL4 has a bimodal color/magnitude plot. Its bimodal distribution is similar to that of RWA6 in the Cygnus OB7 association (Wolk et al. 2013). On the color/magnitude plot in Figure 8 there are two groups of points, both elongated along the extinction direction. The brighter, but redder, group of points is centered at about $J=14.3$ mag and $J - K = 5.05$ mag and fainter but bluer group of points is

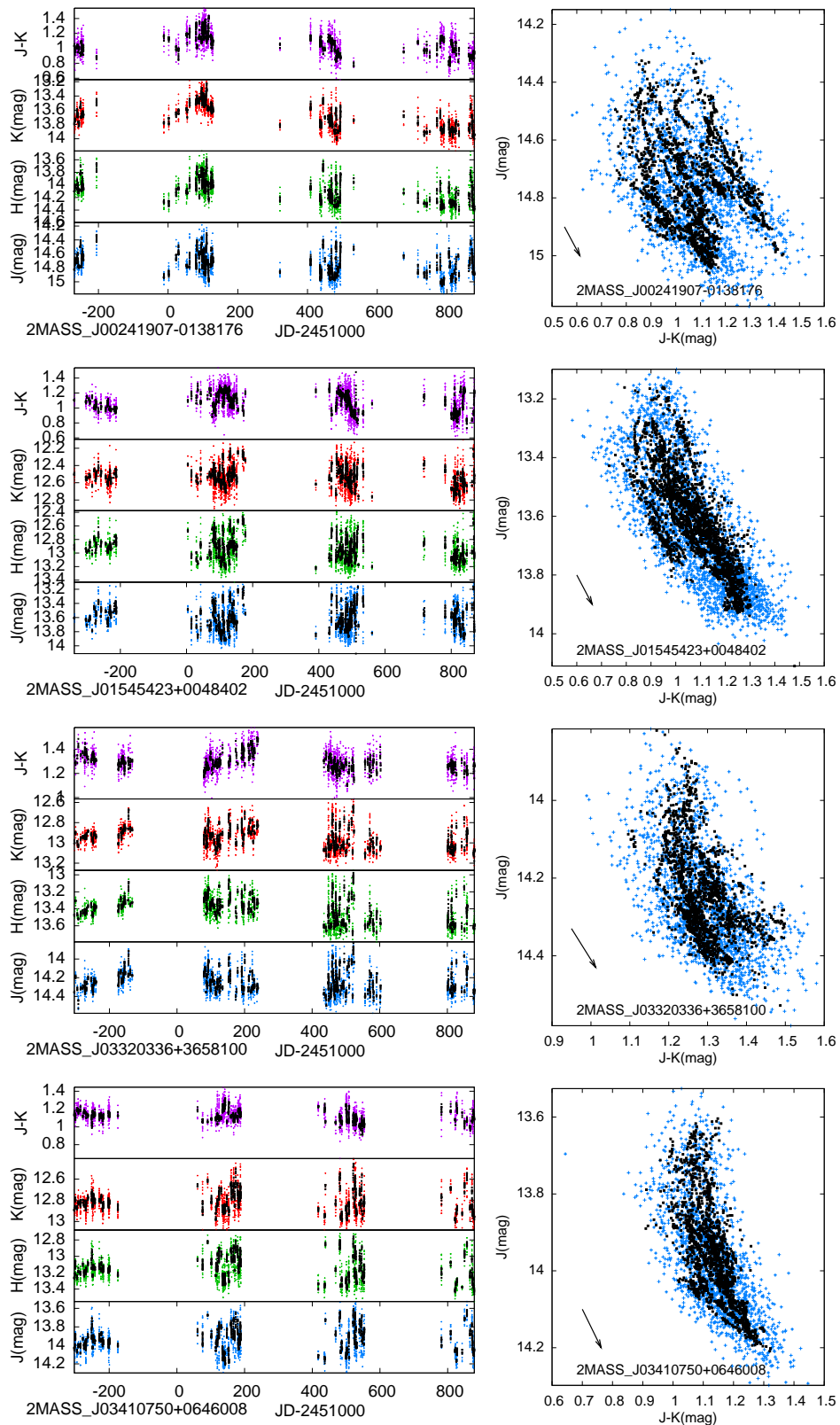


Figure 5. Light curves and colors (similar to figure 3) but of active galaxies found by searching for red dimming events. These become bluer when they are brighter. The objects are listed in Table 4.

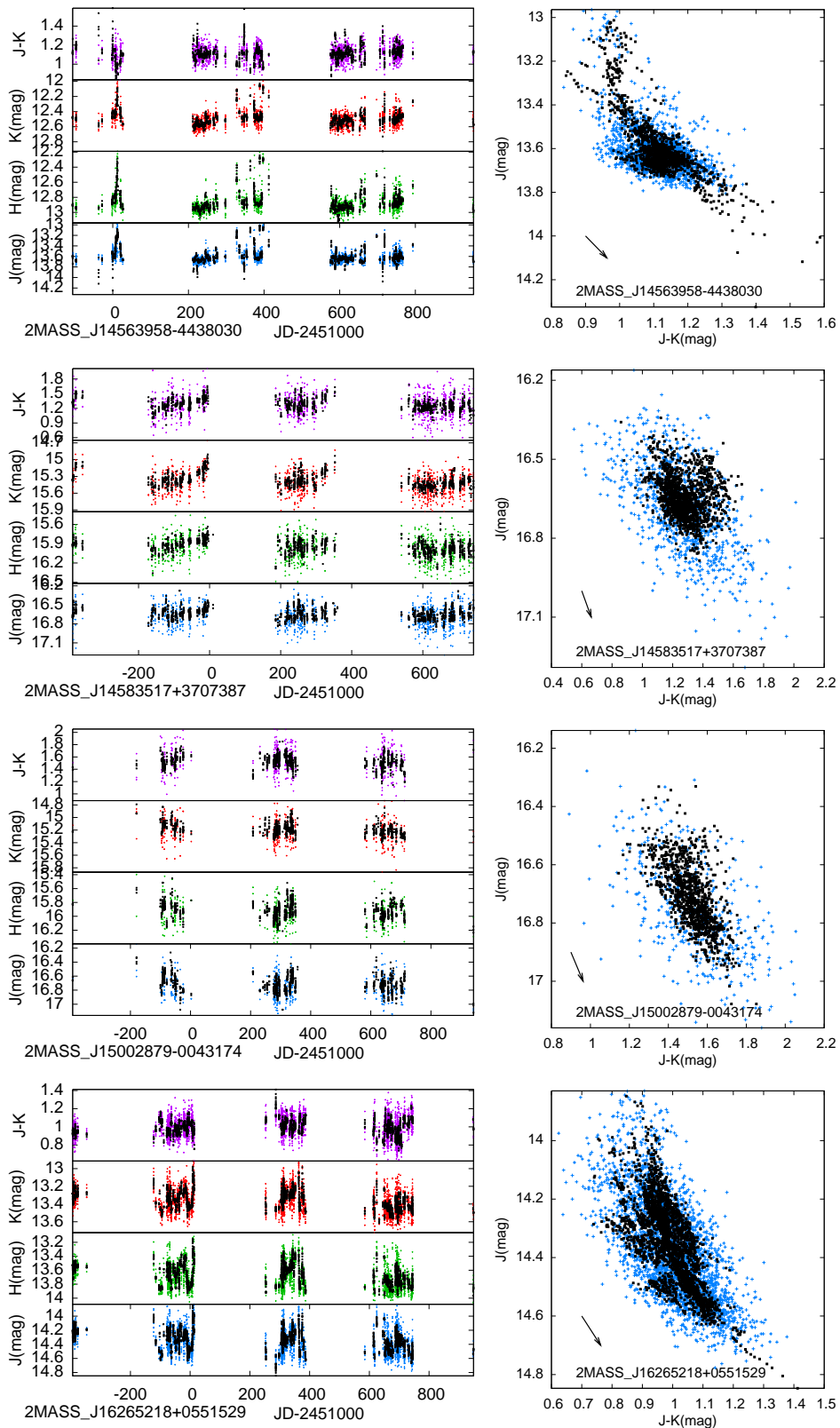


Figure 5. Figure continued.

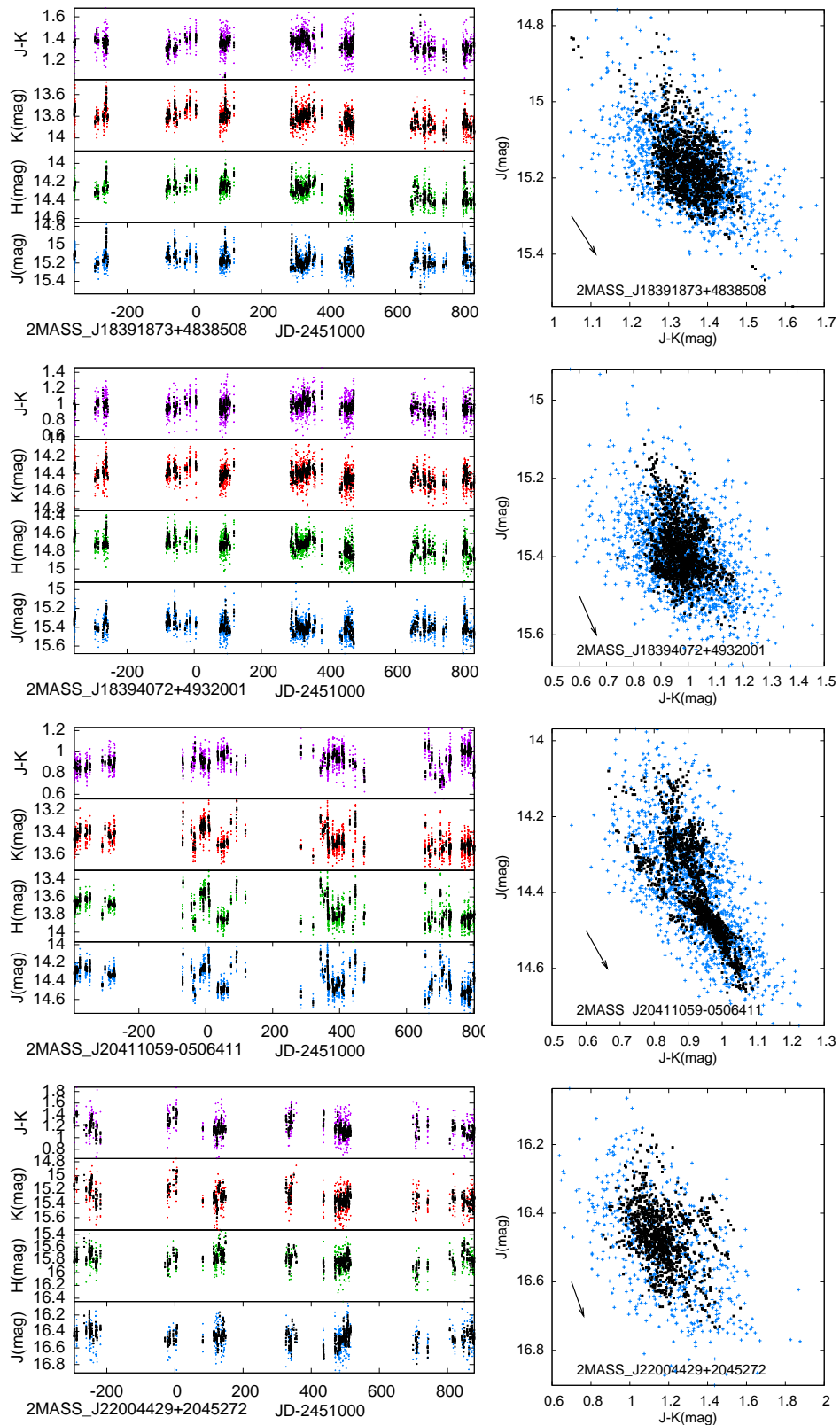


Figure 5. Figure continued.

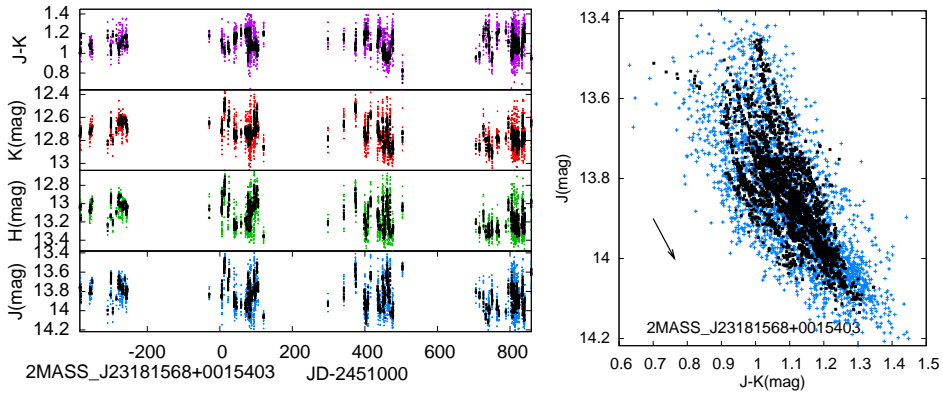


Figure 5. Figure continued.

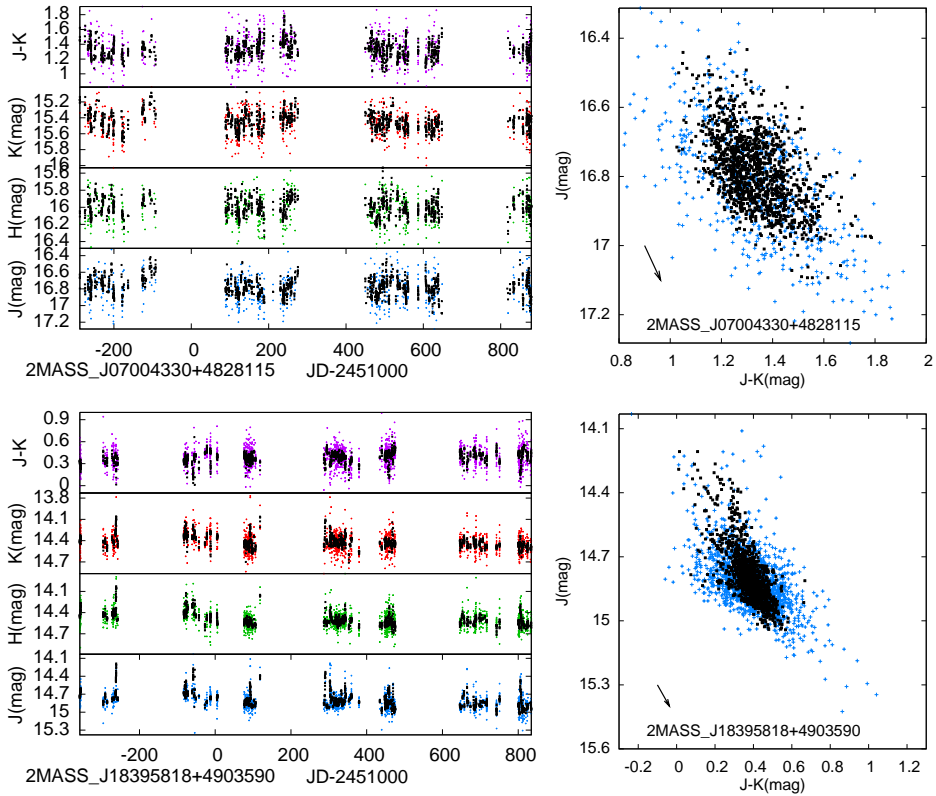


Figure 6. Light curves and colors (similar to figure 3) but of additional objects found in the red dimming search. These objects were not identified as extragalactic sources or young stellar objects and are listed in Table 5.

centered at $J = 14.6$ and $J - K = 4.95$ mag. The object was classified as inverse eclipse/circumbinary disk by Parks et al. (2013) and with a period of 65.61 days, but the period identified as twice as long by Plavchan et al. (2008b) who remarked on the two-state behavior of the object. Plavchan et al. (2008b); Parks et al. (2013) proposed the following model: WL4 is a triple system comprised of three approximately equal mass $0.6 L_{\odot}$ young stellar objects. A close double system, denoted W4a, W4b has a 130.87 or 65.61 day period and the two stellar components WL4a, WL4b, alternate being obscured by a circumbinary disk. The circumbinary disk is tilted with respect to the orbital plane of

the binary due to the third star, W4c, that is resolved by direct imaging. Star spots on the unobscured third object, W4c, account for shorter timescale variability. We ask: is this model consistent with the distribution of points seen on the color/magnitude diagram? The equal number of points in both regions in the color/magnitude diagram are consistent with the fraction of the orbital period spent in each state being large. The only known systems with eclipses lasting a significant fraction of the orbital period are like KH 15D, with a misaligned circumbinary disk. The only addition we would make to this model is that the color magnitude distri-

Table 4. Active galaxy candidates found in a search for red dimming events

Object (1)	J (2)	σ_J (3)	H (4)	σ_H (5)	K (6)	σ_K (7)	$J - K$ (8)	[3.4]-[4.6] (9)	[4.6]-[12] (10)	R (11)	b (12)	Ref (13)
2MASS J00241907-0138176	14.728	0.187	14.066	0.207	13.669	0.186	1.06	0.10	2.40	0.40	0.500±0.026	L11
2MASS J01545423+0048402	13.578	0.223	12.918	0.209	12.532	0.143	1.05	0.26	4.41	0.74	1.078±0.019	H12
2MASS J03320336+3658100	14.225	0.142	13.397	0.166	12.941	0.120	1.28	0.39	3.77	0.51	0.828±0.032	2MES
2MASS J03410750+0646008	13.939	0.162	13.152	0.170	12.818	0.120	1.12	-0.08	2.63	0.64	1.104±0.031	H12
2MASS J08255140-3907590	15.586	0.209	14.386	0.166	13.386	0.110	2.2	1.05 ^a	2.99	0.90	1.402±0.012	-
2MASS J14563958-4438030	13.563	0.244	12.852	0.178	12.485	0.121	1.08	-0.07	0.65	0.48	0.827±0.038	P05
2MASS J14583517+3707387	16.645	0.163	15.940	0.178	15.393	0.187	1.25	0.29	3.50	0.54	0.397±0.018	SDSS
2MASS J15002879-0043174	16.702	0.183	15.914	0.189	15.190	0.174	1.51	0.39	2.87	0.58	0.510±0.026	SDSS
2MASS J16265218+0551529	14.324	0.199	13.652	0.187	13.352	0.132	0.07	-0.08	2.42	0.67	1.167±0.025	P03
2MASS J18391873+4838508	15.162	0.124	14.288	0.113	13.815	0.095	1.35	0.11	2.16	0.50	0.524±0.023	P03
2MASS J18394072+4932001	15.382	0.100	14.729	0.115	14.410	0.120	0.97	0.08	2.68	0.44	0.349±0.017	P03
2MASS J20411059-0506411	14.364	0.143	13.725	0.149	13.460	0.108	0.90	-0.04	2.06	0.66	0.975±0.028	P03
2MASS J22004429+2045272	16.465	0.156	15.788	0.168	15.311	0.175	1.15	0.33	3.54	0.56	0.417±0.022	SDSS
2MASS J23181568+0015403	13.819	0.172	13.115	0.178	12.737	0.117	1.08	0.18	2.92	0.72	1.022±0.019	S10

Here we list extragalactic objects found in a search for red dimming events. Columns: (1): 2MASS point source catalog identifier. (2)-(7): Mean and standard deviations, in magnitudes, computed from the light curves in J , H , and K bands, respectively. (8): $J - K$ Color in magnitudes was computed from the mean magnitudes of the light curve. (9,10): Colors in Vega magnitudes from the WISE survey (Cutri et al. 2012). The four bands of the WISE survey are 3.4, 4.6, 12 and 22 μm . (11): The cross correlation coefficient, R , (equation 1) between color and magnitude. (12): The slope, b , of the best fitting line on the color/magnitude plot (equation 2). (13): Reference where the object was listed as extragalactic. L11 - (Lavaux & Hudson 2011); H12 - (Huchra et al. 2012); 2MES - in the 2MASS Extended Sources catalog; P05 - (Paturel et al. 2005); P03 - (Paturel et al. 2003); SDSS - classified as a galaxy with SDSS imaging; S10 - Schawinski et al. (201). ^a2MASS J08255140-3907590 is identified here as extragalactic based on its red [3.4]-[4.6] color. 2MASS J03320336+3658100 is a variable radio source (Ofek & Frail 2011).

Table 5. Additional objects found in a search for red dimming events

Object (1)	J (2)	σ_J (3)	H (4)	σ_H (5)	K (6)	σ_K (7)	$J - K$ (8)	[3.4]-[4.6] (9)	[4.6]-[12] (10)	R (11)	b (12)
2MASS J07004330+4828115	16.773	0.177	16.003	0.175	15.453	0.170	1.32	0.34	3.03	0.60	0.499±0.021
2MASS J18395818+4903590	14.790	0.169	14.500	0.126	14.411	0.137	0.38	-	-	0.55	0.588±0.022

Here we list objects found in a search for red dimming events that were not identified as extragalactic or young stellar objects. Columns are the same as in Table 4. Objects lacking WISE colors were not in the WISE all sky data release catalog.

butions suggest that both components experience variations in extinction.

Two sources with only moderate value of $J - K$ (~ 3 mag), 2MASS J16273084-2424580 (Iso-Oph 151) and 2MASS J16272738-2431165 (WL 13) display long (two day) colorless dimming events ~ 0.2 mag deep. Iso-Oph 151 is a Class III object, but WL 13 has a similar $J - K$ relatively blue color compared to the other more embedded young stellar objects. Of the young stellar objects found in the dimming search, these two light curves have relatively low values of dispersion in K band (the dispersion in J band in some cases is high because the objects are faint). An obscuring disk need not cause reddening during eclipse if it is opaque or if it is comprised of large dust grains and so has a gray extinction law (e.g., EE-Cep Galan et al. 2012). The long and relatively colorless dimming events seen in these objects may be due to dense or gray material passing in front of the object.

6 SYNTHESIS

In this section we first discuss objects that were found by both searches and then we discuss how different types of ob-

jects might be found or eliminated in automated searches. We then discuss correlation coefficients measured from the color/magnitude distributions for the different classes of objects found in our searches.

6.1 Objects found in both searches

Because our color insensitive dimming search (described in section 3) used the standard deviation of the light curve, variable stars with large standard deviations were often discarded. However, our red dimming event search described below was sensitive to variable stars with large standard deviations.

Which objects were found in both searches? Young stellar objects in the ρ -Ophiucus region that exhibited large and red brightness variations were present in both searches. The object 2MASS J18394250+4856177 and discussed in section 3, was found in both searches. This object exhibits rapid and correlated in color/magnitude variations, similar to the active galaxies, but it was not identified as extragalactic or as having an infrared excess. The long timescale variable star 2MASS J19015057-0431087 was present in both searches. The previously known eclipsing binary 2MASS J08255405-3908441 and the eclipsing bina-

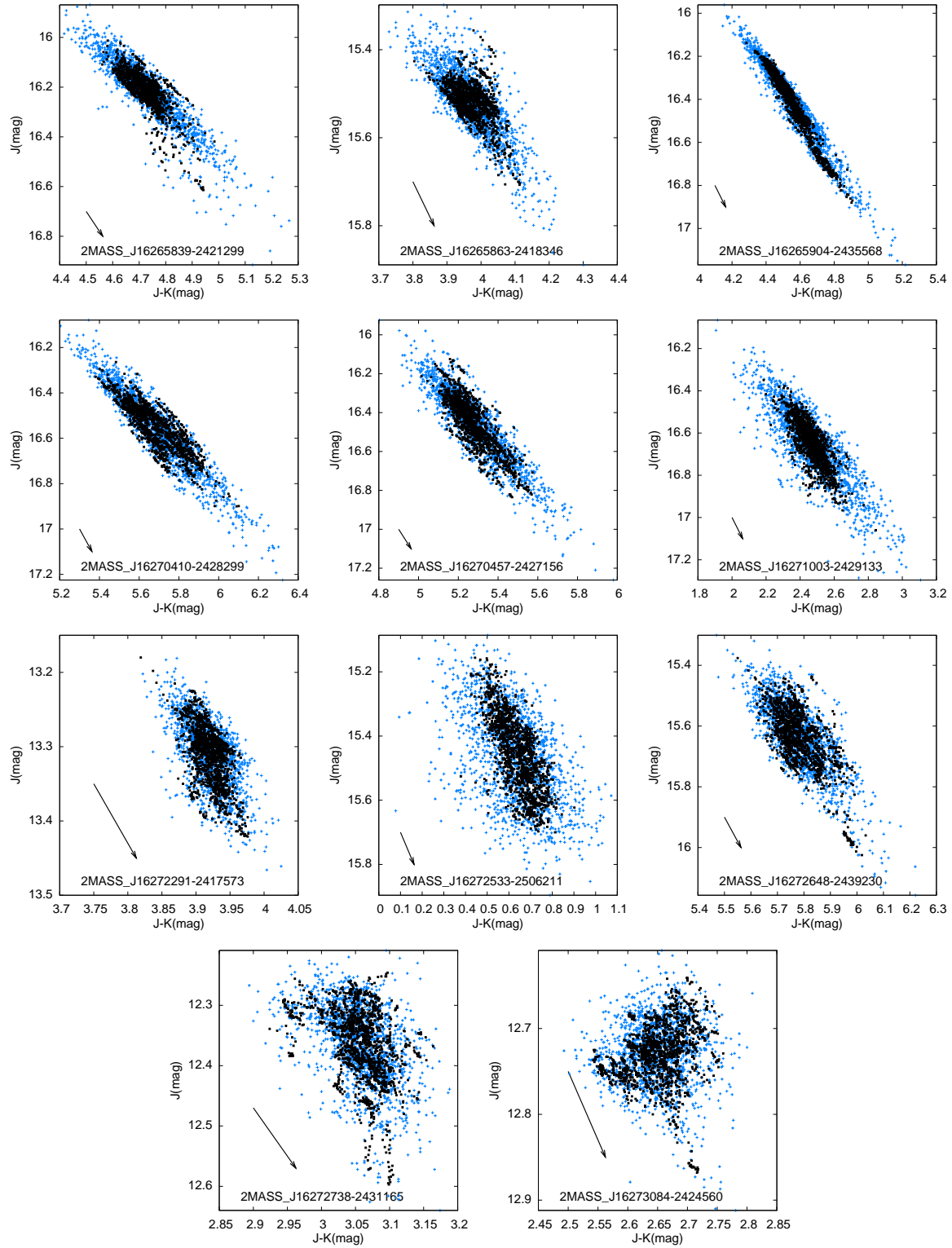


Figure 7. Color/magnitude plots (similar to those in Figure 3) but of young stellar objects in the ρ Ophiucus star formation region found only in the color-blind dimming event search.

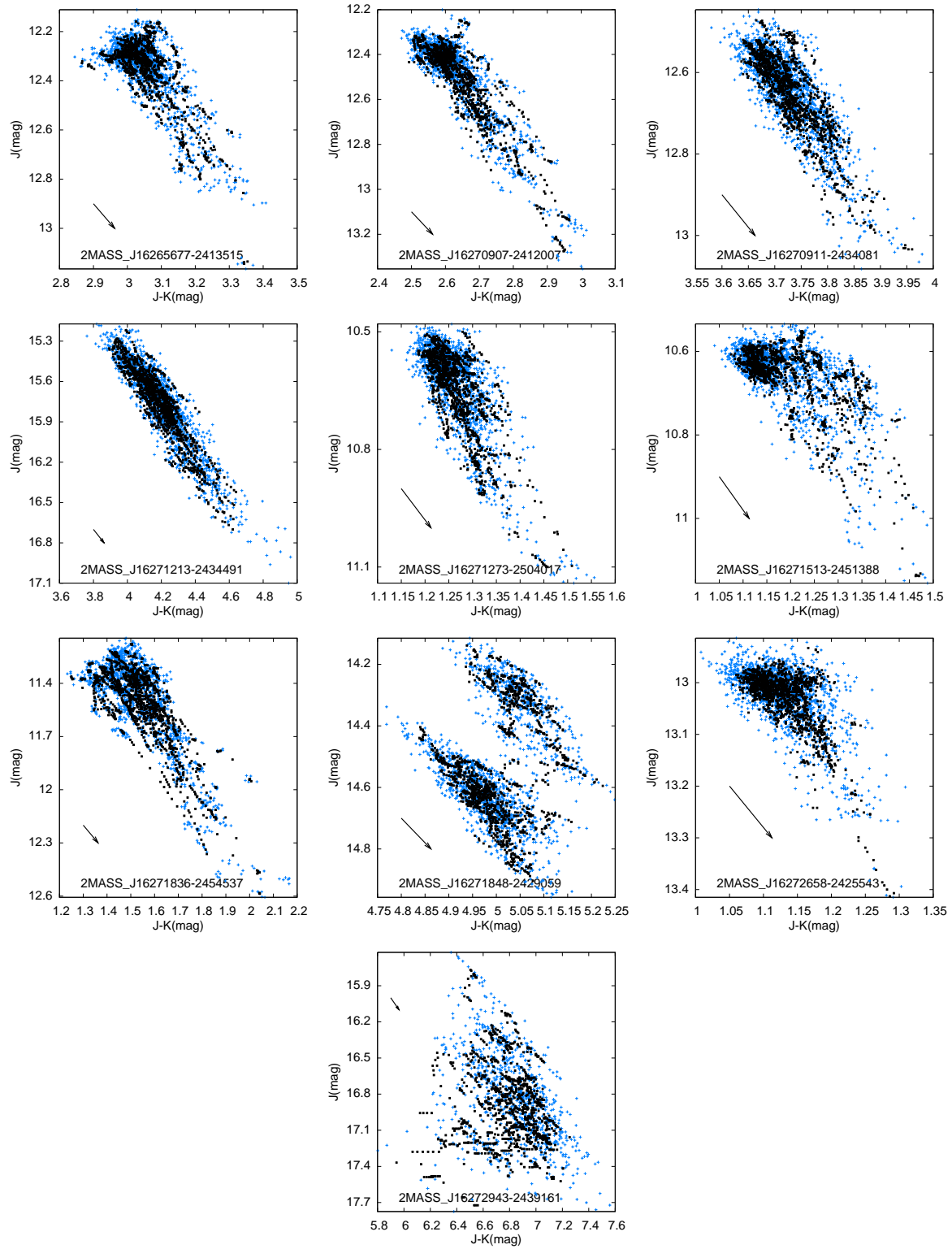


Figure 8. Color/magnitude plots (similar to those in Figure 3) but of young stellar objects in the ρ Ophiucus star formation region found in both dimming event searches. These young stellar objects exhibit deep and red dimming events. Note the bimodal distribution in the color/magnitude plot of WL4 (2MASS J16271848-2429059).

Table 6. Young stellar objects in the ρ Ophiucus region

Object (1)	ID (2)	variability type (3)	J (4)	σ_J (5)	H (6)	σ_H (7)	K (8)	σ_K (9)	R (10)	b (11)
2MASS J16265839–2421299	YLW 1C	P/S/Hot SS	16.213	0.129	13.165	0.068	11.489	0.056	0.91	1.003±0.011
2MASS J16265863–2418346	ISO-Oph 87	Irr/flares	15.519	0.079	12.995	0.048	11.548	0.050	0.80	0.814±0.015
2MASS J16265904–2435568	WL 14	LTV/Ext	16.397	0.175	13.463	0.083	11.867	0.049	0.97	1.130±0.007
2MASS J16270410–2428299	WL 1	LTV	16.539	0.167	13.028	0.061	10.872	0.064	0.94	0.900±0.008
2MASS J16270457–2427156	ISO-Oph 98	P/S/Cool SS	16.445	0.181	13.075	0.108	11.178	0.080	0.91	1.085±0.012
2MASS J16271003–2429133	-	Irr/unknown	16.638	0.175	15.111	0.116	14.179	0.097	0.85	0.878±0.014
2MASS J16272291–2417573	ISO-Oph 135	P/S/Cool SS	13.310	0.046	10.692	0.039	9.390	0.035	0.65	0.953±0.028
2MASS J16272533–2506211	-	P/S/unknown	15.453	0.145	14.985	0.119	14.819	0.143	0.48	0.470±0.022
2MASS J16272648–2439230	YLW 16C	P/S/cool SS	15.628	0.113	12.046	0.078	9.852	0.075	0.78	0.914±0.019
2MASS J16272738–2431165	WL 13	P/S/cool SS	12.365	0.066	10.368	0.066	9.308	0.061	0.45	0.668±0.034
2MASS J16273084–2424560	ISO-Oph 151	LTV/Acc	12.729	0.040	10.998	0.044	10.079	0.063	0.02	0.014±0.021
2MASS J16265677–2413515	ISO-Oph 83	P/S/Ext	12.360	0.148	10.385	0.122	9.302	0.100	0.77	1.395±0.030
2MASS J16270907–2412007	ISO-Oph 106	P/Ecl/Ext	12.501	0.191	10.755	0.150	9.862	0.113	0.90	1.783±0.023
2MASS J16270911–2434081	WL10	P/S/cool SS	12.655	0.106	10.266	0.087	8.923	0.068	0.82	1.514±0.027
2MASS J16271213–2434491	WL11	P/S/hot SS?	15.723	0.292	13.180	0.237	11.573	0.172	0.93	1.732±0.018
2MASS J16271273–2504017	-	P/S/hot starspot?	10.643	0.115	9.763	0.097	9.376	0.082	0.77	1.636±0.035
2MASS J16271513–2451388	-	P/Ecl/Ext	10.659	0.085	9.809	0.070	9.469	0.070	0.65	0.669±0.020
2MASS J16271836–2454537	-	P/S/hot starspot?	11.510	0.267	10.583	0.215	9.966	0.172	0.82	1.526±0.027
2MASS J16271848–2429059	WL4	P/IEcl/CB	14.526	0.181	11.341	0.193	9.528	0.193	0.03	0.078±0.061
2MASS J16272658–2425543	-	P/Ecl/Ext	13.028	0.063	12.368	0.053	11.897	0.049	0.63	0.845±0.026
2MASS J16272943–2439161	YLW 16b	LTV/Ext	16.721	0.395	13.069	0.536	10.249	0.471	0.49	0.818±0.049

The top 11 YSOs were found only in the sigma-limited dimming search and the bottom 10 YSOs found in both searches. Columns: (1): The 2MASS identifier of the star. (2): The identifier used by Parks et al. (2013) in their variability study. (3): The type of variability found by Parks et al. (2013) in their Table 5 for periodic variables, their Table 7 for the aperiodic long timescale variables or in their Table 8 for irregular variables. Stars are either Periodic (P), long-timescale variables (LTV) or irregular (Irr). If periodic, then variability is either sinusoidal (S), eclipse-like (Ecl), or inverse eclipse like (IEcl). Color variations are either similar to hot starspots (hot SS) or cool starspots (cool SS) or due to extinction (Ext). If an LTV then color variations are either colorless (and so associated with accretion; Acc) or red and so due to extinction (Ext). CB stands for circumbinary disk. (4)-(9): Mean and standard deviations, in magnitudes, computed from the light curves in J , H , and K bands, respectively. (10): The cross correlation coefficient, R , (equation 1) between color and magnitude. (11): The slope, b , of the best fitting line on the color/magnitude plot (equation 2).

ries 2MASS J19014028-0424137, 2MASS J19020201-0430160 were also found in both searches. The eclipsing binary, 2MASS J08255405-3908441, has a period of 8.08986 days and was previously identified by Plavchan et al. (2008). They noted the color variations ($\Delta(J-H) \sim 0.1$ mag) in primary eclipses (see their Figure 49) and attributed them to different effective temperatures in the primary and secondary stars. The color magnitude plot for these three eclipsing binaries are shown in Figure 9 and display deviations in color in two directions, one for primary eclipse and the other for the secondary eclipse.

As only two eclipsing binaries, were found in the search for red dimming events, the color variability selection we used in the red dimming search can effectively eliminate most eclipsing binaries. The sigma-limit used in the color-blind search eliminates many, but not all active galaxies and young stellar objects. The red dimming event search did not find a population of objects, with a single or a few rare dimming events, rather the objects found tend to exhibit continuous variability. By combining both searches with an absolute limit on the dispersion in the light curve we could eliminate all the active galaxies and young stellar objects. Such a search on a larger sample of stars might reveal an object such as OGLE-LMC-ECL-11893, with a rare red eclipse.

6.2 Cross correlation coefficients

In Figure 10a we show the slope b fit to the color/magnitude diagrams versus correlation coefficients measured from the color/magnitude distributions. Points plotted are listed in Tables 3, 6, 4, and 5). For objects identified in the red-dimming search, large red diamonds (denoted Red-YSO in the key) are the young stellar objects, large green squares are the extragalactic objects (Red-GAL in the key), turquoise triangles are additional objects not identified as extragalactic (Red-Add in the key). For objects identified in the color insensitive, sigma-limited dimming search small red diamond are the young stellar objects only found in the color insensitive search (CB-YSO in the key), extragalactic objects are small green squares, (CB-Gal in the key), blue stars are aperiodic objects (CB) and the large red stars are eclipsing binaries, 2MASS J08255405-3908441, 2MASS J19014028-0424137, 2MASS J19020201-0430160 (denoted RED-EB in the key) found in both searches. Using the same symbols and objects we plot the dispersion in the H band light curve, σ_H , versus the cross correlation coefficient R in Figure 10b. We have used the H -band dispersion as some of the young stellar objects were faint in J band.

Points on the right of Figure 10a,b have high correlation coefficients and so when they dim, they become redder. Points on the left have uncorrelated variations in color and brightness. Points on the top of 10a have large changes in brightness for a small change in color. A slope of 1.6 is consis-

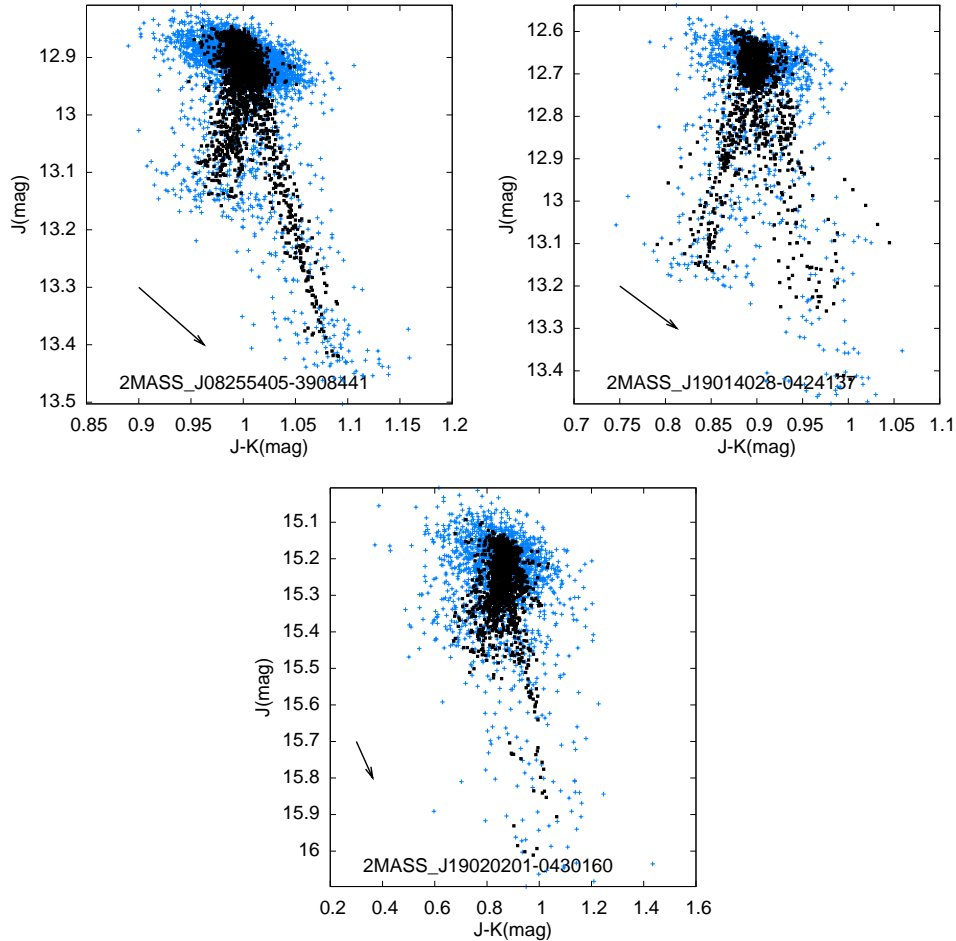


Figure 9. Color/magnitude plot for 2MASS J08255405-3908441, and 2MASS J19014028-0424137 and 2MASS J19020201-0430160, the eclipsing binaries with color variations in eclipse (plots are similar to those in figure 3). These eclipsing binaries were found in both red-dimming and sigma-limited colorblind dimming searches.

tent with variability caused by a variation in extinction. The young stellar objects have the largest slopes and correlation coefficients. The galaxies are intermediate in both correlation coefficient and slope. The intermediate value of slope is perhaps expected as the color variations are expected to be due to a contribution from a blue active source. Objects exhibiting low correlation coefficient and slope are most likely eclipsing binaries but we have failed to find their periods, however, this group could contain more exotic objects. The odd eclipsing binary with red eclipses stands out as having a high slope but a low correlation coefficient.

Two young stellar objects have a low correlation coefficients, WL4 (2MASS J16271848-2429059) that has a bimodal color/magnitude plot (see that in Figure 8) and Iso-Oph 151 (2MASS J16273084-2424560) displaying a significant, but rare dip in brightness.

The AGNs and young stellar objects tend to lie on different regions in Figure 10a, however, they are not completely separated. However, the AGNs tend to have higher dispersions in their light curve, as seen from Figure 10b. It is likely that most of the AGNs could be eliminated by placing a limit on the magnitude dispersion σ_H .

7 SUMMARY AND CONCLUSION

We have presented a search for dimming events in approximately 40,000 objects (mostly stars) observed over a period of about 4 years in the 2MASS Calibration fields. We searched for day long dimming events that are significant compared to the standard deviation of the object's magnitude distribution. Our goal was to search for rare long dimming by an exotic object such as the occulting disk of OGLE-LMC-ECL-11893, consequently we did not restrict our search to objects with periodic variability.

In a color independent search for day-long dimming events we found 46 previously unidentified eclipsing binaries and 6 additional periodic variable stars. We also found a few objects exhibiting dimming events but without strong color variations during dimming. These could be eclipsing binaries but we have not observed enough eclipses to identify their period, or they could be something more exotic. We also found two active galaxies and a few objects displaying correlated color and magnitude variations suggesting that they are also extragalactic. Young stellar objects in the ρ Ophiucus field, previously studied by Parks et al. (2013) also displayed dimming events, however no ordinarily non-variable or quiescent star in this region was detected with a

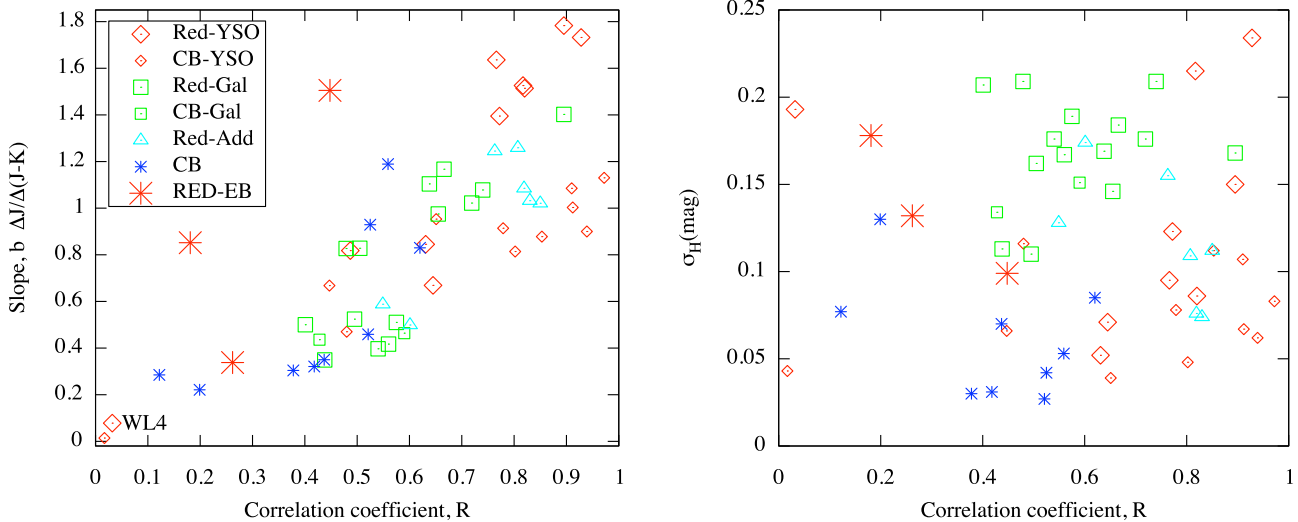


Figure 10. a) Slope on color/magnitude plot, b vs cross correlation coefficient, R for objects listed in Tables 3, 6, 4, and 5). For objects identified in the red-dimming search, large red diamonds (denoted Red-YSO in the key) are the young stellar objects, large green squares are the extragalactic objects (Red-GAL in the key), turquoise triangles are additional dimming objects not identified them as extragalactic (Red-Add in the key) For objects identified in the color insensitive, sigma-limited dimming search small red diamond are the young stellar objects only found in the color insensitive search (CB-YSO in the key), extragalactic objects are small green squares, (CB-Gal in the key), blue stars are aperiodic objects (CB) and the large red stars are the eclipsing binaries, (denoted RED-EB in the key) found in both searches. A slope of 1.6 is consistent with variability caused by changes in extinction. The young stellar objects have the largest slopes and correlation coefficients. The galaxies are intermediate and the objects found in the color independent search have low correlation coefficients. b) The same objects but plotting the standard deviation σ_H in H band of the light curve versus the cross correlation coefficient R .

dimming event. Two less obscured objects (Iso-Oph 151, a Class III source and WL 13 a Class II source) variable young stellar objects in this field displayed 0.2 mag deep, colorless two-day long dimming events. We failed to find any ordinarily quiescent object with a dimming event lasting more than a day.

From the 4 year survey of approximately 40,000 objects (and excluding the young stellar objects) approximately 1 of 700 is an eclipsing binary, and approximately 1 of 4000 exhibits a dimming event but requires additional observations to classify.

We carried out a second search, looking for dimming events that are accompanied by reddening. In this search we did not restrict the search by the standard deviation of the object’s magnitude distribution. This search primarily revealed two classes of objects, active galaxies that become bluer as they increase in brightness, and Class II young stellar objects in the ρ Ophiucus region that exhibit large variations in brightness (up to a magnitude in J band) that are also red ($\Delta(J - K) \sim 0.4$ mag). The young stellar objects found in both searches display color variations consistent with interstellar extinction, though in some cases there must be additional emission mechanisms influencing the brightness and color. Of the 40000 objects searched we found 14 active galaxies, corresponding to approximately one of 3000 searched, exhibiting dimming events and correlations between brightness and color.

We did not find any dimming events greater than 0.2 magnitude deep and longer than a day outside the ρ Ophiucus region. We did not find an object similar to OGLE-LMC-ECL-11893, ordinarily without variability, but exhibiting

long, deep and red dimming events, in the 2MASS calibration fields (and including objects in the ρ Ophiucus field).

Because we failed to find any ordinarily quiescent object with a long (greater than a day) dimming event, we consider how we would eliminate more common objects in a future larger scale search. The red dimming event search effectively eliminates all but three eclipsing binaries comprised of two stars with significantly different effective temperatures. The only targets that were found in both searches were the young stellar objects in the ρ Ophiucus star forming region, the eclipsing binaries with color variations, a long period variable and a rapidly varying object that might be an active galaxy.

For objects found in our searches we computed correlation coefficients between color and magnitude and fit a slope to points in the color magnitude distribution. We find that young stellar objects, active galaxies and objects displaying insignificant variations in color (and so are most likely eclipsing binaries) tend to be found in different regions of this plot, however these two quantities are not sufficient to definitely separate the populations. Most active galaxies can be eliminated with a limit on the dispersion in the magnitude distribution. Perhaps combining additional color information or by using higher moments of the magnitude distributions, or additional spectral information the two populations could be better delineated and identified.

We note that our search criteria were not standard compared to those used in other surveys, so our search for variable stars in the 2MASS calibration fields is not complete. For example, we did not follow up bright stars identified with

large dispersions computed from the magnitude distribution in the light curve.

There is a long history of searching for and following up objects that temporarily brighten, such as novae or gravitational micro-lenses, but there have been few systematic searches for dimming events. In this study we took care to discard multiple dimming events that occurred at the same time in the same calibration field. However, differential photometric methods, where the brightness of nearby stars are contrasted in the data reduction pipeline, can be used to minimize the number of spurious dimming events (e.g., see discussion by van Eyken et al. 2011). This would allow a more robust search for rare and single dimming events in a larger sample of stars and make it possible to develop strategies to follow up unidentified objects.

New and ongoing surveys are increasing the number of stars photometrically monitored, improving upon the cadence of observations and on the numbers of photometric bands for photometric monitoring. An object similar to OGLE-LMC-ECL-11893 could be found in larger and multi-color survey that searched for deep and red dimming events, a high cadence single band survey that examined the shape of the light curve during a single dimming event or computed the magnitude distribution in an eclipse (Meng et al. 2014), or an extended (many year) lower cadence single band survey survey that examined the shape of the eclipse in the period-folded light curve (Dong et al. 2014; and how it was found). Rare events, such as a transit of a planet hosting icy rings, like Saturn, or a disk with large dust grains (and a grayer extinction law) might not cause a red dimming event but could induce an atypically shaped eclipse. Differentiating this from an eclipsing binary would require high cadence and high photometric precision.

We searched here for rare dimming events using a multicolor near-infrared survey. An optical survey could survey a larger number of fainter stars, however a near-infrared survey would be particularly useful for studying star forming regions embedded in the Galactic plane. Eclipsing disks associated with planet formation might better be found in a near-infrared survey that includes many young stars, whereas a search for ringed planets might be better done in the visible bands.

In our search of approximately 40000 light curves approximately 1 of 4000 objects (11 objects total) that have displayed dimming events remain unidentified. As more common objects such as eclipsing binaries and active galaxies are classified in a continuously monitored field, it could become possible to discover more exotic objects displaying rare dimming events. Roughly scaling from our estimates, a search for dimming events in a billion good (few thousand points spanning a few years) light curves would yield of order a million eclipsing binaries, a few hundred thousand active galaxies and a few hundred thousand unidentified objects (that would include young stellar objects). Amongst these could be a bright counterpart to OGLE-LMC-ECL-11893 or an extrasolar analog of Saturn's rings occulting its host star.

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